



HAL
open science

Infrared study of transitional disks in Ophiuchus with Herschel

Isabel Rebullido, Bruno Merín, Álvaro Ribas, Ignacio Bustamante, Hervé Bouy, Pablo Riviere-Marichalar, Timo Prusti, Göran L. Pilbratt, Philippe André, Péter Ábrahám

► **To cite this version:**

Isabel Rebullido, Bruno Merín, Álvaro Ribas, Ignacio Bustamante, Hervé Bouy, et al.. Infrared study of transitional disks in Ophiuchus with Herschel. *Astronomy and Astrophysics - A&A*, 2015, 581, pp.A30. 10.1051/0004-6361/201425556 . cea-01383675

HAL Id: cea-01383675

<https://cea.hal.science/cea-01383675>

Submitted on 19 Oct 2016

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

Infrared study of transitional disks in Ophiuchus with *Herschel*^{★,★★,★★★,★★★★}

Isabel Rebullido¹, Bruno Merín¹, Álvaro Ribas^{1,2,3}, Ignacio Bustamante^{1,2,3}, Hervé Bouy², Pablo Riviere-Marichalar¹,
Timo Prusti⁴, Göran L. Pilbratt⁴, Philippe André⁵, and Péter Ábrahám⁶

¹ European Space Astronomy Centre (ESA), PO Box, 78, 28691 Villanueva de la Cañada, Madrid, Spain
e-mail: Isabel.Rebullido@sciops.esa.int

² Centro de Astrobiología, INTA-CSIC, PO Box – Apdo. de correos 78, 28691 Villanueva de la Cañada, Madrid, Spain

³ ISDEFE – ESAC, PO Box, 78, 28691 Villanueva de la Cañada, Madrid, Spain

⁴ ESA, Scientific Support Office, Directorate of Science and Robotic Exploration, European Space Research and Technology Centre (ESTEC/SRE-S), Keplerlaan 1, 2201 AZ Noordwijk, The Netherlands

⁵ Laboratoire AIM Paris, Saclay, CEA/DSM, CNRS, Université Paris Diderot, IRFU, Service d'Astrophysique, Centre d'Études de Saclay, Orme des Merisiers, 91191 Gif-sur-Yvette, France

⁶ Konkoly Observatory, Research Centre for Astronomy and Earth Sciences, Hungarian Academy of Sciences, PO Box 67, 1525 Budapest, Hungary

Received 19 December 2014 / Accepted 2 July 2015

ABSTRACT

Context. Observations of nearby star-forming regions with the *Herschel* Space Observatory complement our view of the protoplanetary disks in Ophiuchus with information about the outer disks.

Aims. The main goal of this project is to provide new far-infrared fluxes for the known disks in the core region of Ophiuchus and to identify potential transitional disks using data from *Herschel*.

Methods. We obtained PACS and SPIRE photometry of previously spectroscopically confirmed young stellar objects (YSO) in the region and analysed their spectral energy distributions.

Results. From an initial sample of 261 objects with spectral types in Ophiuchus, we detect 49 disks in at least one *Herschel* band. We provide new far-infrared fluxes for these objects. One of them is clearly a new transitional disk candidate.

Conclusions. The data from *Herschel* Space Observatory provides fluxes that complement previous infrared data and that we use to identify a new transitional disk candidate.

Key words. stars: formation – stars: pre-main sequence – protoplanetary disks – planets and satellites: formation

1. Introduction

Protoplanetary disks around young stars are objects of major interest as they lead us to a better understanding of star and planet formation. Transitional disks are key in this study since they appear to have an unusual radial structure. They have been proposed as the environment for planet formation (Marsh & Mahoney 1992) and have other proposed formation mechanisms, such as photo-evaporation by ultraviolet light emitted by the central star (Clarke et al. 2001), grain growth (Dullemond & Dominik 2005), and gravitational instabilities (see Espaillat et al. 2014, for a recent review on transitional disks). Recently,

Kim et al. (2013) studied accretion towards transitional disks in Orion A and concluded that planet formation was the most likely explanation for their observations. What characterizes a transitional disk is a lack of excess in the near- or mid-IR region (usually around 8 or 10 μm) and typical Class-II excesses in mid- to far-IR. This lack of near- and mid-IR excess denotes an inner disk opacity hole, which is related to the dust distribution in the surroundings of the star and reveals inner holes. These objects are thought to be an intermediate stage between Class II objects (optically thick disks) and Class III objects (smaller amount of material in the disk).

The first transitional disks were reported by Strom et al. (1989). Since that time, the known population of these objects has grown substantially thanks to data from *Spitzer* Space Telescope (Werner et al. 2004). The study of these objects has improved with new and more powerful telescopes, such as the *Herschel* Space Observatory (Pilbratt et al. 2010), which provides a wider range of wavelengths. *Herschel* also represents an improvement *Spitzer*'s sensitivity and spatial resolution at long wavelengths, which allows for a reduction in the noise level of measurements at these wavelengths. Studies in the millimetric and sub-millimetric range allowed for the imaging and direct measurement of hole sizes, such as those made

* *Herschel* is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.

** Final reduced *Herschel* maps are only available at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/qcat?J/A+A/581/A30>

*** Appendix A is available in electronic form at <http://www.aanda.org>

**** All tables are also available at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/qcat?J/A+A/581/A30>

Table 1. Sensitivity of the *Herschel* observations used in this study.

	F_{70} [mJy]	F_{160} [mJy]	F_{250} [mJy]	F_{350} [mJy]	F_{500} [mJy]
Sensitivity per scan	12.2	14.3	12.6	10.5	15.0
Effective sensitivity	2	2.4	6.3	5.2	7.5
Minimum flux detected	20	50	70	10	30

by [Andrews et al. \(2011\)](#) with the Sub-Millimeter Array ([Ho et al. 2004](#)). Recent work with ALMA ([Wootten & Thompson 2009](#)) has achieved new results in the field as seen in works by [van der Marel et al. \(2013\)](#) and the recent study on the transitional disk HL Tauri ([ALMA Partnership et al. 2015](#)).

The aim of this work is to study the young stellar objects (YSOs) in the Ophiuchus star-forming region. *Herschel* data provide us with accurate fluxes of the detected objects, and enables the construction of the spectral energy distributions (SEDs) along with other multi-wavelength photometric data, collected by [Ribas et al. \(2014\)](#). The study of the SEDs also allows us to classify the transitional disks in the region by identifying the previously mentioned lack of excess in near and/or mid-IR region.

The structure of this work is as follows: Sect. 2 describes the *Herschel* observations and data reduction, explaining the source detection and photometry extraction processes. In Sect. 3 we explain the results obtained from the data reduction, both with the method described in [Ribas et al. \(2013\)](#) and with the analysis of the SED of each individual object. In Sect. 4 we discuss the detection statistics and compare them with other similar studies. In Sect. 5 we present the conclusions of this work.

2. Observations and data reduction

The Ophiuchus cloud complex was observed by *Herschel* as part of the *Herschel* Gould Belt Survey ([André et al. 2010](#)) and by a deeper PACS survey ([Alves de Oliveira et al. 2013](#)). It is one of the closest star-forming regions, located at an estimated distance of 130 pc and with an age between 2 and 5 Myr, although it has been suggested that it is younger (see [Wilking et al. 2008](#), for a review on the region). Apart from its proximity, most significant characteristic of this cloud is its dense core, which is the focus of our study. The cloud itself appears as a large scale structure with complex filaments, heated by B-type stars, and is strongly emitting at infrared wavelengths. We might consider the cloud itself as a possible source of contamination. Figure 1 shows the region with the objects in the sample overplotted and marked differently according to their state of non-detected, detected, or transitional.

The maps used to examine this region were obtained from two sets of observations, using the parallel mode of both PACS ([Poglitsch et al. 2010](#)) and SPIRE ([Griffin et al. 2010](#)) on board the *Herschel* Space Observatory. For SPIRE maps (250 μm , 350 μm and 500 μm) the parallel mode was used, with a speed of 60''/s (program *KPGT_pandre_1*), but for the PACS maps (70 μm and 160 μm) the scan mapping mode was used, with a cross-scan speed of 20''/s (program *OT1_pabraham_3*). We used this later PACS data because it goes deeper, as shown by the sensitivity limits given in the SPIRE/PACS Observers' Manual and in Table 1 in this work. In both cases (PACS and SPIRE), a single pair of scan and cross-scan was obtained. The obsids for the SPIRE maps are (1342-) 205093 and 205094, and for PACS maps (1342-) 238816 and 238817.

The maps were produced using Scanamorphos ([Roussel 2013](#)), an IDL software designed to process *Herschel* maps. Version 24.0 was used for PACS maps, with calibration file

version 65, and version 22.0 for SPIRE maps, with calibration files version 12.3.

2.1. Sample selection and point source photometry

The objective of this work is to determine the nature of the sources detected in the maps produced with PACS and SPIRE observations and to obtain fluxes in the mid- to far-IR regions when detected. The YSO candidates were selected from the work of [Ribas et al. \(2014\)](#) (which for Ophiuchus contains objects from [Natta et al. 2002](#); [Wilking et al. 2005](#); [Alves de Oliveira et al. 2010](#); [Erickson et al. 2011](#)), for having known spectral type from spectroscopy, making a list of 258 sources, all of them located in the core of the cloud. For a more complete study, the objects from [Cieza et al. \(2010\)](#) contained in the region of interest were also considered. The final sample consists of 261 objects. We used photometry from the 2MASS *K*-band ([Skrutskie et al. 2006](#)) and WISE4 ([Wright et al. 2010](#)) to classify the disks according to the method described by [Lada \(1987\)](#). The only object that does not fulfill the criteria is IRS 48, which has a positive slope that is almost flat. This object cannot be classified as Class I and remains Class II in the final classification because of the lack of strong emission in mid- and far-IR.

To determine whether our candidates are transitional disks or not, the first step was to detect sources in our maps and extract the photometry from them. We used the algorithm *Sussextractor* in the *Herschel* interactive processing environment (HIPE), version 12.1, with a threshold of $S/N > 3$. Additionally, visual inspection was applied to all sources. Only sources clearly separated from filaments and distinguishable from the background were selected as valid detections.

We extracted aperture photometry for these sources using the sourceExtractorDaophot task in HIPE ([Ott 2010](#)). The fluxes were aperture corrected to account for the shape of the PSF and they are listed in Table 4. Those corrections were obtained from [Balog et al. \(2014\)](#) for PACS and from the SPIRE Observers' Manual (now named SPIRE Handbook) version 2.5. In Table 4, objects marked with an asterisk suffered from high nearby background emission and, in those cases, we used a special sky estimate from a rectangular area identified after visual inspection. The set of apertures and the corrections applied to the photometry extracted are shown in Table 2. We tested several point-source extraction algorithms, including *Sussextractor*, *Daophot*, *AnnularSkyAperturePhotometry*, and *Hyper* ([Traficante et al. 2015](#)), and different apertures. The combination above provides the best fit to the MIPS70 ([Rieke et al. 2004](#)) fluxes of clean selected sources in the field.

At the end of the process, we had 49 successfully detected sources in at least one PACS band, and 19 in a SPIRE band; PACS also detected all of these. Images of the detected objects can be found in Fig. A.2, for visual inspection.

The calibration errors for PACS and SPIRE are 5% and 7% (as stated in the PACS Observer's Manual, version 2.3 and in the SPIRE Observer's Manual version 2.5), however, as a more conservative estimation, we used 25% for both. These measurement uncertainties account for the high variations in the

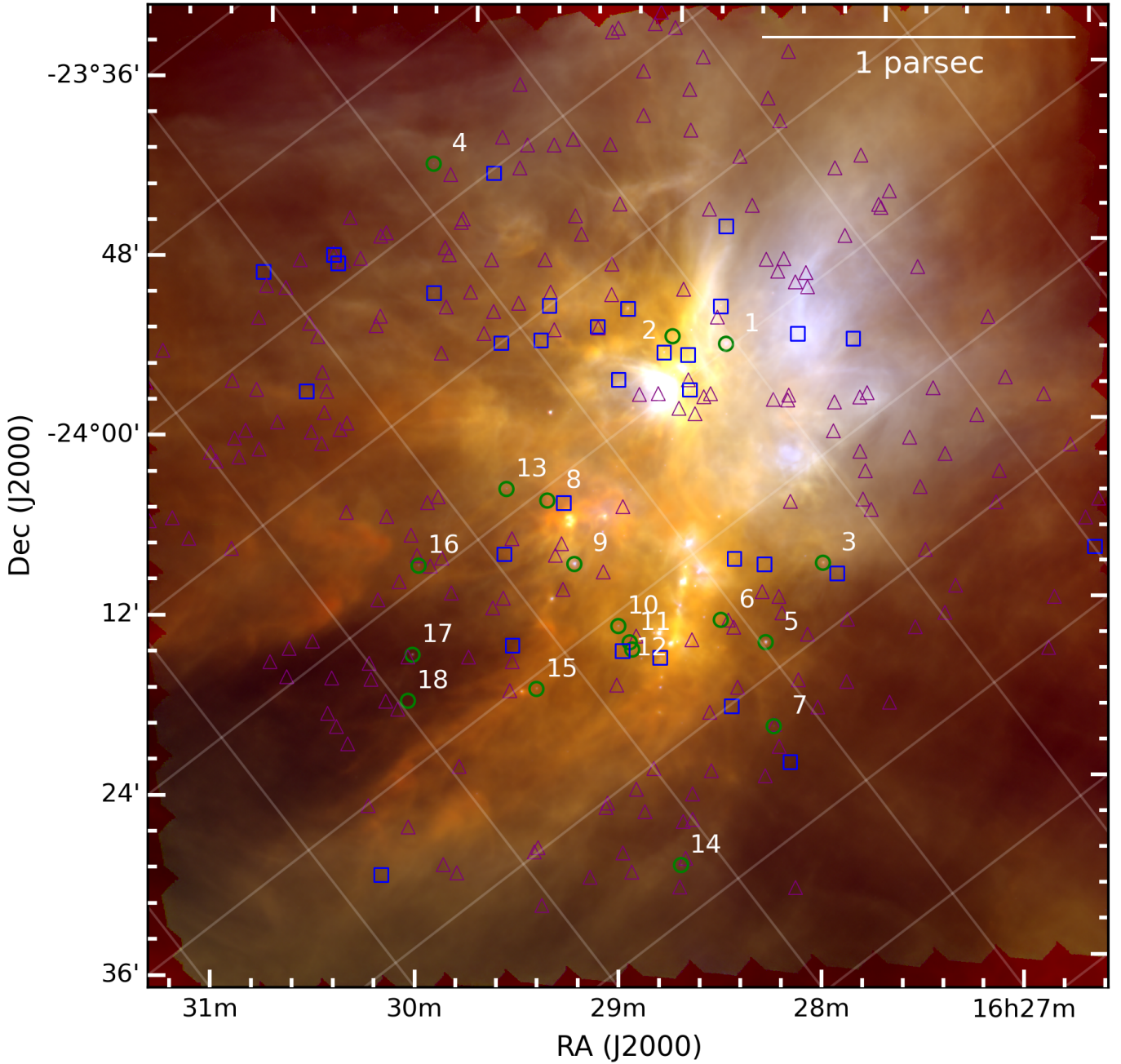


Fig. 1. RGB image (PACS 70 in blue, PACS 100 in green, and PACS 160 in red) of the observed Ophiuchus region with the marked objects. Transitional disks are marked with green circles. The blue squares mark the detected objects, and the purple triangles mark the rest of the sample of known young stars in the region.

results of the different extraction methods and apertures listed above.

In Table 3 we show the parameters of the detected sources, as given in the literature, and for an easier reference we list the identification numbers (ID) used in this work as seen in Fig. 1. From now on, the ID number is referenced in parenthesis for its respective source. Table 4 gives the fluxes measured with *Herschel* for each source.

2.2. Sensitivity and non-detections

The high background present in the region increases the sensitivity limit of these observations, and affects the number of

detections. In Table 1 we present the sensitivities given in the SPIRE/PACS Parallel Mode Observers' Manual v2.1 (Sect. 2.3), and the minimum flux detected for each band. Also, the effective sensitivity is calculated taking the number of scans made per map into account; this includes six for the PACS maps and two for the SPIRE maps. We also report a lower value of fluxes in the cleanest areas of the map, which agrees with the fact that the extended emission from the molecular cloud reduces the effective sensitivity.

Another study took into account the sources detected by *Spitzer* and catalogued by c2d (Evans et al. 2009) in each of the apertures defined per band and per detected source to estimate possible contamination due to *Herschel* resolution. The flux in

Table 2. Aperture photometry parameters.

Band	<i>FWHM</i> (")	Radius (")	Inner annulus (")	Outer annulus (")	Aperture correction factors
PACS-70	5.4	6	25	35	1.5711
PACS-160	10.5	12	25	35	1.4850
SPIRE-250	18	22	60	90	1.2584
SPIRE-350	24	30	60	90	1.2242
SPIRE-500	36	42	60	90	1.1975

Table 3. Parameters of detected sources as extracted from the literature.

ID.	Name	RA (deg)	Dec (deg)	SpT	A_v^*	References
–	2MASS J16245974-2456008	16:24:59.63	–24:55:59.32	M3.5	0.10	1
–	V*V852Oph	16:25:24.38	–24:29:43.77	M4.5	0.21	2
–	2MASS J16253673-2415424	16:25:36.73	–24:15:42.32	K4	0.10	1
–	2MASS J16253958-2426349	16:25:39.58	–24:26:34.27	M2	0.13	2
–	V*V2058Oph	16:25:56.18	–24:20:47.77	K4.5	0.61	1
1	Haro1-6	16:26:03.02	–24:23:36.04	K1	5.70	4
–	2MASS J16261684-2422231	16:26:16.83	–24:22:23.32	K6	0.11	1
2	DoAr 24	16:26:17.08	–24:20:21.47	K4.5	0.13	2
–	2MASS J16262189-2444397	16:26:21.88	–24:44:39.67	M8	0.10	2
–	DoAr 24E	16:26:23.28	–24:20:59.37	G6	0.10	2
3	DoAr 25	16:26:23.68	–24:43:13.57	K5	0.21	2
–	GSS32	16:26:24.03	–24:24:48.32	K5	0.63	1
–	2MASS J16262407-2416134	16:26:24.08	–24:16:13.27	K5.5	0.19	2
–	2MASS J16263297-2400168	16:26:32.98	–24:00:16.77	M4.5	0.09	2
–	2MASS J16263682-2415518	16:26:36.93	–24:15:52.32	M0	0.55	1
–	[GY92]93	16:26:41.28	–24:40:17.87	M5	0.09	2
–	2MASS J16264285-2420299	16:26:42.83	–24:20:30.32	M1	0.11	1
–	2MASS J16264643-2412000	16:26:46.48	–24:11:59.97	G3.5	0.10	2
4	WSB 40	16:26:48.58	–23:56:34.57	K5.5	0.09	2
–	WL18	16:26:48.98	–24:38:25.07	K6.5	0.59	2
–	2MASS J16265677-2413515	16:26:56.68	–24:13:51.47	K7	0.11	2
5	SR 24 S	16:26:58.48	–24:45:36.67	K1	0.18	2
6	2MASS J16270659-2441488	16:27:06.68	–24:41:49.07	M5.5	0.09	2
–	2MASS J16270907-2412007	16:27:09.03	–24:12:01.32	M2.5	0.11	1
7	WSB 46	16:27:15.08	–24:51:38.77	M2	0.09	2
–	[WMR2005]4 – 10	16:27:17.48	–24:05:13.67	M3.5	0.10	2
–	2MASS J16271836-2454537	16:27:18.38	–24:54:52.77	M3.75	0.16	2
–	WSB 49	16:27:22.98	–24:48:07.07	M4.25	0.09	2
–	2MASS J16272658-2425543	16:27:26.58	–24:25:54.47	M8	0.14	2
8	2MASS J16273084-2424560	16:27:30.88	–24:24:56.37	M3.25	0.12	2
–	2MASS J16273311-2441152	16:27:33.13	–24:41:14.32	K6	0.50	1
9	IRS 48	16:27:37.23	–24:30:34.32	A0	0.76	1
10	IRS 49	16:27:38.28	–24:36:58.67	K5.5	0.15	2
–	2MASS J16273832-2357324	16:27:38.28	–23:57:32.97	K6	0.14	2
11	2MASS J16273863-2438391	16:27:38.60	–24:38:39.00	M6	0.24	3
–	2MASS J16273901-2358187	16:27:38.98	–23:58:19.17	K5.5	0.15	2
12	WSB 52	16:27:39.48	–24:39:15.87	K5	0.09	2
13	SR 9	16:27:40.28	–24:22:04.37	K5	0.17	2
–	2MASS J16274270-2438506	16:27:42.68	–24:38:51.27	M2	0.11	2
–	V*V2059Oph	16:27:55.58	–24:26:18.27	M2	0.11	2
–	2MASS J16280256-2355035	16:28:02.58	–23:55:03.61	M3	4.30	4
–	2MASS J16281379-2432494	16:28:13.83	–24:32:49.32	M4	0.10	1
14	2MASS J16281385-2456113	16:28:13.83	–24:56:10.32	M0	0.10	1
15	WSB 60	16:28:16.58	–24:36:58.57	M4.5	0.11	2
–	2MASS J16281673-2405142	16:28:16.83	–24:05:15.32	K6	0.08	1
16	SR 20W	16:28:23.38	–24:22:40.87	K5	0.50	2
17	SR 13	16:28:45.28	–24:28:19.27	M3.75	0.20	2
18	2MASS J16285694-2431096	16:28:57.03	–24:31:09.32	M5.5	0.13	1
–	2MASS J16294427-2441218	16:29:44.28	–24:41:21.80	M4	0.80	4

Notes. Also, we give the ID number used to identify the transitional disks. ^(*) All A_v are calculated according to Ribas et al. (2014).

References. 1) Erickson et al. (2011); 2) Wilking et al. (2005); 3) Natta et al. (2002); 4) Cieza et al. (2010).

Table 4. Point source fluxes of each of the 49 sources detected.

ID	Name	F_{70} [Jy]	F_{160} [Jy]	F_{250} [Jy]	F_{350} [Jy]	F_{500} [Jy]
–	2MASS J16245974-2456008	0.05 ± 0.01	–	–	–	–
–	V*V852Oph	0.73 ± 0.18	–	–	–	–
–	2MASS J16253673-2415424	0.74 ± 0.18	–	–	–	–
–	2MASS J16253958-2426349	0.92 ± 0.23	–	–	–	–
–	V*V2058Oph	3.86 ± 0.97	3.77 ± 0.94	12.56 ± 3.14	–	–
1	Haro1-6	10.70 ± 2.68	7.48 ± 1.87	–	–	–
–	2MASS J16261684-2422231	0.20 ± 0.05	–	–	–	–
2	DoAr 24	0.50 ± 0.12	–	–	–	–
–	2MASS J16262189-2444397	0.10 ± 0.02	–	–	–	–
–	DoAr 24E	4.17 ± 1.04	3.91 ± 0.98	–	–	–
3	DoAr 25	1.39 ± 0.35	3.52 ± 0.88	4.17 ± 1.04	5.30 ± 1.33	2.05 ± 0.51
–	GSS32	3.70 ± 0.93	–	–	–	–
–	2MASS J16262407-2416134	3.04 ± 0.76	6.10 ± 1.53	4.73 ± 1.18	2.66 ± 0.66	1.12 ± 0.28
–	2MASS J16263297-2400168	0.08 ± 0.02	0.13 ± 0.03	–	0.01 ± 0.00	–
–	2MASS J16263682-2415518*	1.37 ± 0.34	0.72 ± 0.18	0.44 ± 0.11	–	–
–	[GY92]93	0.02 ± 0.00	–	–	–	–
–	2MASS J16264285-2420299	0.65 ± 0.16	–	–	–	–
–	2MASS J16264643-2412000	0.57 ± 0.14	–	–	–	–
4	WSB 40	0.42 ± 0.10	0.29 ± 0.07	–	–	–
–	WL18	0.42 ± 0.10	–	–	–	–
–	2MASS J16265677-2413515	0.21 ± 0.05	0.79 ± 0.20	–	–	–
5	SR42S	9.70 ± 2.42	7.88 ± 1.97	5.06 ± 1.26	3.02 ± 0.76	1.55 ± 0.39
6	2MASS J16270659-2441488	0.10 ± 0.02	–	–	–	–
–	2MASS J16270907-2412007	0.07 ± 0.02	0.23 ± 0.06	–	0.11 ± 0.03	–
7	WSB 46	0.25 ± 0.06	0.15 ± 0.04	–	–	–
–	[WMR2005]4-10	0.18 ± 0.04	0.32 ± 0.08	0.51 ± 0.13	0.37 ± 0.09	–
–	2MASS J16271836-2454537*	0.07 ± 0.02	0.05 ± 0.01	0.07 ± 0.02	–	–
–	WSB 49	0.04 ± 0.01	–	–	–	–
–	2MASS J16272658-2425543	0.04 ± 0.01	–	–	–	–
8	2MASS J16273084-2424560	0.50 ± 0.13	–	–	–	–
–	2MASS J16273311-2441152	1.27 ± 0.32	0.08 ± 0.02	–	–	–
9	IRS48	37.57 ± 9.39	12.85 ± 3.21	6.48 ± 1.62	2.51 ± 0.63	–
10	IRS49*	1.29 ± 0.32	0.95 ± 0.24	0.07 ± 0.02	0.67 ± 0.17	–
–	2MASS J16273832-2357324*	0.63 ± 0.16	0.25 ± 0.06	0.52 ± 0.13	0.30 ± 0.07	–
11	2MASS J16273863-2438391	0.09 ± 0.02	–	–	–	–
–	2MASS J16273901-2358187	0.44 ± 0.11	0.52 ± 0.13	0.10 ± 0.02	0.09 ± 0.02	0.03 ± 0.01
12	WSB 52*	2.37 ± 0.59	2.23 ± 0.56	3.43 ± 0.86	0.76 ± 0.19	0.28 ± 0.07
13	SR 9	0.86 ± 0.21	0.22 ± 0.05	–	–	–
–	2MASS J16274270-2438506	0.02 ± 0.01	–	–	–	–
–	V*V2059Oph	0.06 ± 0.01	–	–	–	–
–	2MASS J16280256-2355035	0.02 ± 0.01	–	–	–	–
–	2MASS J16281379-2432494	0.05 ± 0.01	–	–	–	–
14	2MASS J16281385-2456113	0.74 ± 0.18	0.68 ± 0.17	0.29 ± 0.07	0.19 ± 0.05	0.14 ± 0.03
15	WSB 60	0.87 ± 0.22	1.19 ± 0.30	1.26 ± 0.32	1.79 ± 0.45	–
–	2MASS J16281673-2405142	0.12 ± 0.03	–	–	–	–
16	SR 20W	0.89 ± 0.22	0.60 ± 0.15	0.57 ± 0.14	0.46 ± 0.12	0.37 ± 0.09
17	SR 13	1.36 ± 0.34	1.32 ± 0.33	0.79 ± 0.20	0.51 ± 0.13	0.20 ± 0.05
18	2MASS J16285694-2431096	0.02 ± 0.01	–	–	–	–
–	2MASS J16294427-2441218	0.05 ± 0.01	–	–	–	–

Notes. (*) Sources near bright background emission; sky measurement was done in clean regions.

MIPS-24 band of each of this c2d sources was used to extrapolate via the median SED of the Ophiuchus region (see Table 5) the expected fluxes at 70 μ m. In Table A.1, we list a sum of all the expected fluxes for this contaminating sources, where it is possible to check whether the expected contamination is lower than the flux error. Notice that this contamination flux is only an estimation, as only the reported object has been detected in the aperture.

3. Results

The fluxes of all the detected sources per band are given in Table 4, as mentioned above. As expected, the higher flux corresponds to the most known object in the region, IRS 48 (#9). All the objects proposed as transitional disk candidates present relatively high fluxes, being 2MASS J16285694-2431096 (#18) the candidate with the lower flux at 70 μ m (20 mJy). The detection

on more than one band for these objects seems to be arbitrary. Less than half of the candidates have been detected either in one band or in all of them. This is probably related to the physical properties of each of the disks rather than to instrumental issues.

3.1. Identification of transitional disks

Once we measured the fluxes, we use the criterion by Ribas et al. (2013) to classify transitional disks with *Herschel* photometry. As previously noted, transitional disks display little or no excess at near- to mid-IR, and large excess at long wavelengths, which translates into a change of slope in the SED of the object around $12\ \mu\text{m}$, i.e. from a negative to positive slope. To proceed with the identification, we defined two indexes, following:

$$\alpha_{\lambda_1-\lambda_2} = \frac{\log(\lambda_1 F_{\lambda_1}) - \log(\lambda_2 F_{\lambda_2})}{\log(\lambda_1) - \log(\lambda_2)}, \quad (1)$$

where λ is measured in μm and F_λ in $\text{erg s}^{-1} \text{cm}^{-2}$. The first index is defined for the band K acquired from 2MASS and $12\ \mu\text{m}$ acquired from WISE. The second index is defined for the $12\ \mu\text{m}$ band and the $70\ \mu\text{m}$ band acquired from *Herschel*. The criterion to determine whether an object is a transitional disk candidate is $\alpha_{12-70} > 0$, since we define transitional disks as objects with a deficit of excess flux at near- to mid-IR fluxes, but standard excesses at longer wavelengths. For DoAr 24 (#2), the $12\ \mu\text{m}$ -band was not available and, therefore, we used $8\ \mu\text{m}$ band from *Spitzer*/IRAC (Fazio et al. 2004) instead.

To understand the result of this process, we represent in Fig. 2 a slope-slope diagram, which shows the value of the two different slopes in the two axes. The figure shows that we obtained two candidates with a positive slope between $12\ \mu\text{m}$ and $70\ \mu\text{m}$, clearly separated from the Class II population. These two candidates, 2MASS J16281385-2456113 (#14) and Haro1-6 (#1), clearly fulfil the criteria of having a lack of mid-IR excess since Haro1-6 (#1) is an object previously classified as debris disk in Cieza et al. (2010). The candidate 2MASS J16281385-2456113 (#14), has never been reported before as a transitional disk. The other objects above the threshold do not have this kind of clear $12\ \mu\text{m}$ flux deficit, which is indicative of the presence of a flatter slope in the mid- to far-infrared wavelengths. If we construct the SEDs of those objects (see Fig. 4), we see that in general they do not present the characteristic gap expected from a transitional disk. However, as they fulfil the criteria, we classified them here as tentative candidates. Some of these candidates were also classified as transitional disks in previous works. Object WSB 60 (#15) was imaged in Andrews et al. (2009), detecting a small inner hole in dust continuum observations done with the Sub-Millimeter Array (Ho et al. 2004). In the work by Cieza et al. (2010) DoAr 25 (#3) was already suggested as a candidate. We consider objects SR 42 S (#5), WSB 46 (#7), and SR 20 W (#16) as tentative candidates despite they do not fulfill the criteria for their nominal values, but we are considering a high error in PACS photometry. In particular, SR 42 S (#5) has already been classified as transitional disk, and appears in Espaillat et al. (2014) as so. For WSB 46 (#7) and SR 20 W (#16), further study is needed to determine their nature.

3.2. Complementary identification with spectral energy distributions

To better analyse the nature of these objects, we built their SEDs using data from optical to mid-infrared from both ground-based and space telescopes (all references for the photometry can be

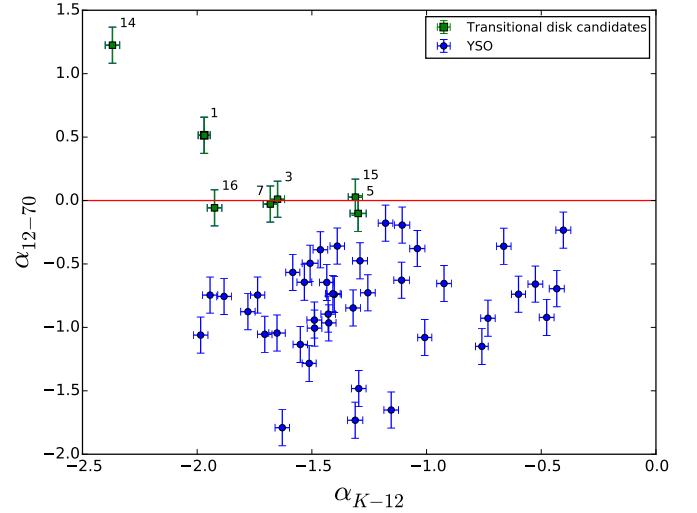


Fig. 2. SED slope between 12 and $70\ \mu\text{m}$ as a function of the SED slope between the *K*-band and $12\ \mu\text{m}$. Transitional disks are marked with green squares.

Table 5. Median SED of detected disks in Ophiuchus.

Band	Median	First quartile	Third quartile	Detections
<i>J</i>	1.0000	1.0000	1.0000	49
<i>H</i>	0.6109	0.5664	0.6577	49
<i>K</i>	0.3305	0.3018	0.3662	49
IRAC-3.6	0.0743	0.0628	0.0958	45
IRAC-4.5	0.0380	0.0305	0.0537	47
IRAC-5.8	0.0212	0.0162	0.0293	46
IRAC-8.0	0.0121	0.0088	0.0190	48
MIPS-24	0.0018	0.0013	0.0030	48
PACS-70	0.0003	0.0001	0.0008	49
PACS-160	8.53×10^{-5}	5.54×10^{-5}	0.0002	26
SPIRE-250	3.48×10^{-5}	2.32×10^{-5}	0.0001	17
SPIRE-350	9.60×10^{-6}	5.88×10^{-6}	1.44×10^{-5}	15
SPIRE-500	2.62×10^{-6}	2.12×10^{-6}	4.18×10^{-6}	8

found in Sect. 2.1). We completed these SEDs with the photometry shown in Table 2 and extracted from *Herschel* data, therefore, we cover a range between 0.35 to $500\ \mu\text{m}$. Figures 4 and A.1 show the obtained SEDs, plus the NextGen atmosphere model for each object (Hauschildt et al. 1999), which are the best approximation of how a naked photosphere would emit as a function of its spectral type.

We also built the median SED of all the objects detected in the region for comparison (see Table 5), which is plotted along each object's SEDs. Because of the lack of detections for fluxes under the sensitivity limits given in Table 4, the median SED might be slightly overestimated, but we assume the effect in our result is negligible.

To determine the interstellar extinction, we used the procedure in Ribas et al. (2014) and the extinction law from Weingartner & Draine (2001), which uses a model of grains to estimate the interstellar extinction, scattering, and infrared emission. In each plot of Figs. 4 and A.1, the observed fluxes are also shown as empty circles. The de-reddened fluxes, according to the associated interstellar extinction, are shown as filled circles.

When visually inspecting the SEDs, we detect objects with a lack of mid-IR excess that were not classified as transitional disks, according to the criteria of Ribas et al. (2013) described in Sect. 3.1. Some of these objects are well known in the literature, such as IRS 48 (#9). To make a more reliable study of the region,

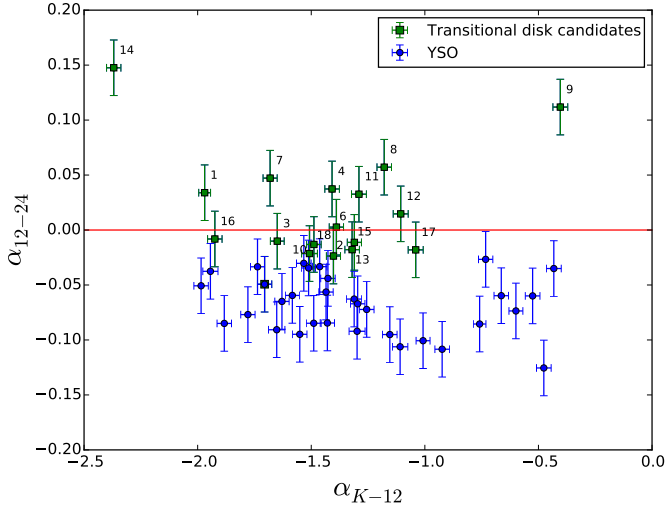


Fig. 3. SED slope between 12 and 24 μm as a function of the SED slope between the K -band and 12 μm . Transitional disks are marked with green squares.

we add in this subsection a complementary criterion based on *Spitzer* data to identify the other transitional disks or transitional disks candidates detected by *Herschel* with a change of slope between 12 μm and 24 μm . Figure 3 shows the new slope-slope diagram.

We detected 17 transitional disks candidates with this criteria. Six of the objects detected with the previous criteria appear now as well, with the exception of SR 24 S (#5), which is a confirmed transitional disk.

Figure 4 shows the SEDs for the 18 transitional disks candidates, where the change of slope is not equally noticeable for all of them. In the case of 2MASS J16281385-2456113 (#14), the change of slope is evident, in agreement with that expected from the slope-slope diagram. In cases such as WSB 60 (#15) or IRS 48 (#9), which are well-known transitional disks, the slope is flatter, and located in a different position in the wavelength axis. For SR 24 S (#5), we see an unexpected behaviour, where we get a discrepancy between MIPS-24 and WISE4 fluxes, but we still see an increase in the far-IR emission with respect to the near-IR. Another object classified as a transitional disk in Cieza et al. (2010) is SR 9 (#13), although it has a continuous decreasing slope, as the criterion shows, and looks more like a Class II SED. The rest of the objects in the sample have been widely studied, but not previously considered transitional disks. The variation in the outputs of both methods illustrates the complexity in defining a selection criteria.

4. Discussion

4.1. Detection statistics

The initial sample was composed of 261 YSO objects in the centre of Ophiuchus, with known spectral type from optical spectroscopy and all classified as Class II objects. Our sample is different from that in Evans et al. (2009), since they had photometrically selected objects, including many objects with other classes, while we have only spectroscopically confirmed Class II YSOs. All of these YSOs fell within the coverage of the maps used, and 49 were detected in at least one *Herschel* band, 49 in PACS, and 19 in SPIRE. This leads to a *Herschel* detection rate of $18.77\% \pm 2.6\%$, which is much smaller than the percentage of detections obtained in similar studies in

Chamaeleon (Ribas et al. 2013) and Lupus (Bustamante et al. 2015) of around 30%. Given that Ophiuchus is closer than those regions (150 ~ 200 pc), the low detection rate is probably due to the higher background, which is emitting at mid- and long-IR wavelengths, and precludes the detection of faint objects.

4.2. Incidence of transitional disks in the centre of Ophiuchus

We report here the detection of 18 transitional disk candidates in the cloud complex in the centre of Ophiuchus based on new *Herschel* and previous known data of spectroscopically confirmed YSO sample. Despite the fact that all of these objects fulfil either one or both of the criteria exposed previously, only a few of them have evident changes of slopes in the SEDs. The candidate with the biggest change in slope, 2MASS J16281385-2456113 (#14), is new to the literature.

The fraction of transitional disk candidates observed in Ophiuchus based on our *Herschel* sample, is $37\%_{-6}^{+7}$. We have considered two classification criteria, depending on the change of slope, that is, on the wavelength position of the lack of mid-IR excess. If we only consider our main criteria, the fraction is lower, at $14\%_{-4}^{+6}$, but it is compatible with the fractions measured in other regions with similar ages in previous works (Espaillat et al. 2014). Even though we have identified 18 transitional disks candidates in total, some of them present a relatively flat slope and, in the slope-slope diagram, are represented very close to, or even below, the threshold. Having objects with nearly a Class-II slope explains the high fraction of transitional disks in the region, as many of them might be fulfilling the criteria due to the large error in their PACS fluxes. These objects would need further study for their safe classification. One of the objects close to the threshold, however, namely YLW 58 or WSB 60 (#15), had been imaged with SMA and shows a small but conspicuous inner hole (Andrews et al. 2011). Despite this, most of the objects in the first criterion were also detected by the complementary criterion; SR 42 S (#5) was not, even though it is a confirmed transitional disk (Andrews et al. 2011). The fact that *Herschel* data is including transitional candidates to the sample, confirms the improvement that *Herschel* represents regarding reaching to further regions in disks. The longer wavelengths now accessible allow us to detect wider and larger cavities, which we could not have identified solely with data from *Spitzer* at mid-infrared ranges.

4.3. Other interesting objects

Even though we have combined two criteria to identify new candidates, and that these criteria select even very small changes of slope, there might be objects in the sample with the characteristics of a transitional disk, which are, as previously noted, a lack of excess around 10 μm and normal excesses at mid- to far-IR. If we inspect Fig. A.1, we observe that several objects present these features. There are two clear cases of objects previously classified as transitional disk candidates: 2MASS J16280256-2355035 and 2MASS J16294427-2441218. Both of these candidates present a lack of excess around 8 μm and excesses in longer wavelengths, but because of the restrictions in the criteria, neither appear as transitional candidates. The objects V* V852 Oph, 2MASS J16253958-2426349 and 2MASS J16262189-2444397 are similar cases that were never classified as transitional disks, but have a gap in their SED, which could be an indication of a hole in their disks. The object 2MASS J16272658-2425543 also shows a lack of excess around 10 μm with a larger excess in

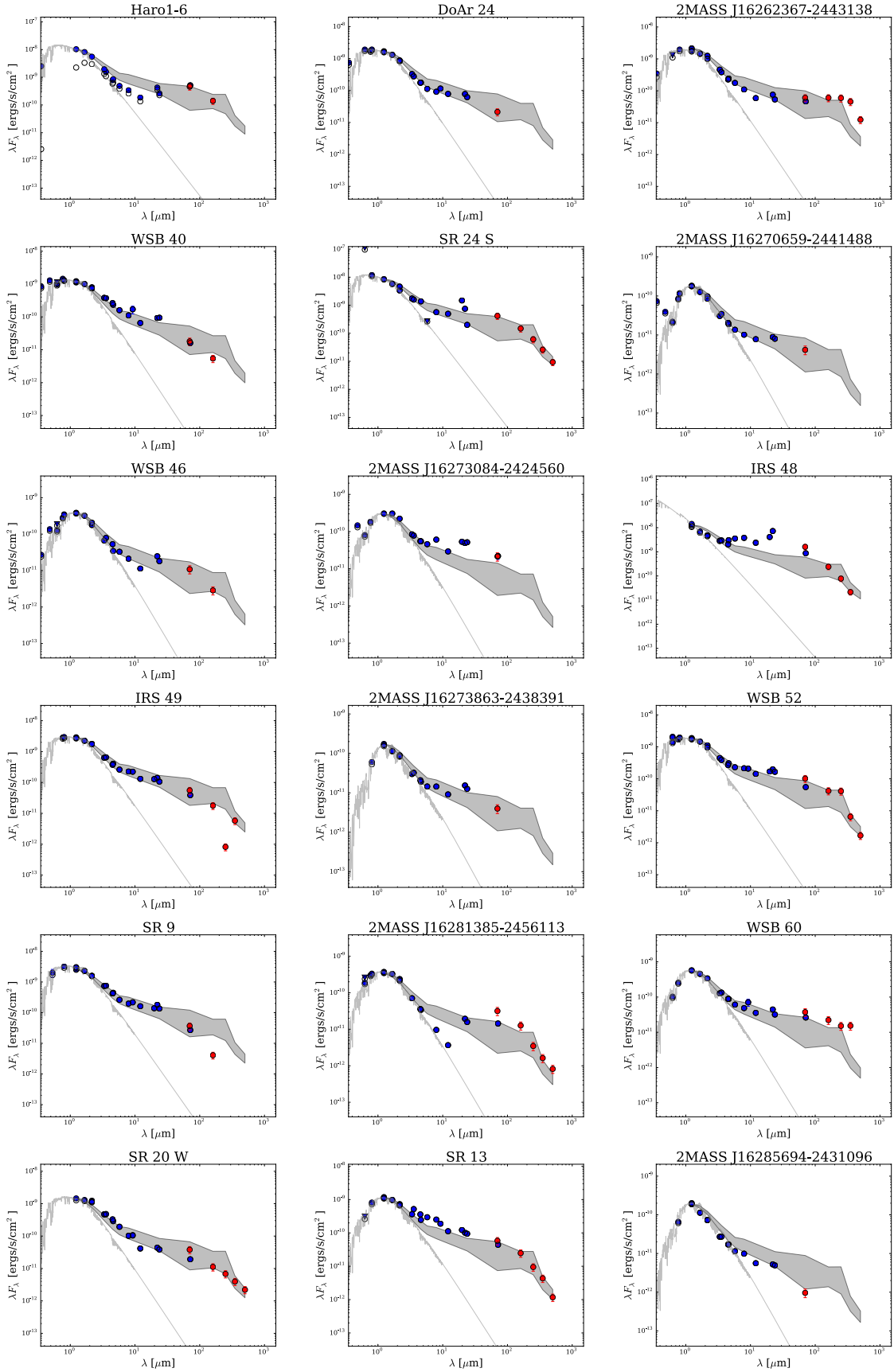


Fig. 4. Spectral energy distribution (SED) of the sources classified as transitional disks candidates. Blue dots show data acquired from the literature, red dots are photometric fluxes obtained from *Herschel* data. Grey dashed line is the photosphere model according to the spectral type, and the grey shaded area is the filled area between the first and third quartile of all the disk fluxes. Observed fluxes are shown with empty circles and A_v values used are in Table 3.

Table 6. Summary of all transitional disk candidates in Ophiuchus.

ID	Name	Classification criteria
1	Haro1-6	12–70*
2	DoAr24	12–24
3	DoAr25	12–70*
4	WSB40	12–24
5	SR42S	12–70
6	2MASS J16270659-2441488	12–24
7	WSB46	12–70*
8	2MASS J16273084-2424560	12–24
9	IRS48	12–24
10	IRS49	12–24
11	2MASS J16273863-2438391	12–24
12	WSB52	12–24
13	SR9	12–24
14	2MASS J16281385-2456113 [†]	12–70*
15	WSB60	12–70*
16	SR20W	12–70*
17	SR13	12–24
18	2MASS J16285694-2431096	12–24
–	V* V852Oph	SED
–	2MASS J16253958-2426349	SED
–	2MASS J16262189-2444397	SED
–	2MASS J16272658-2425543	SED
–	2MASS J16280256-2355035	SED
–	2MASS J16294427-2441218	SED

Notes. (*) These objects have also been classified with the complementary criterion of 12–24. ([†]) This transitional disk candidate is new to the literature.

longer wavelengths. This object could also be considered a transitional candidate, despite the fact that it does not fulfil any of the criteria. All of these objects are shown in Table 6 along with the previous objects classified as transitional disk candidates.

4.4. Study of 70 micron fluxes in transitional disks

For a comparison of the 70 μm flux of both the transitional disk candidates and tentative candidates, we constructed the median SED of all the detected objects in the sample, using the photometric data in Table 5. This would show if, apart from a lack of excess in the near mid-IR, transitional disks also show another remarkable characteristic that could be used in classification or, possibly, in disk modelling. The median SED lacks the contribution from the faintest objects, which might remain undetected by *Herschel* or some of the other surveys, and hence is just an upper limit to the true median SED of Ophiuchus.

When we overplot the median SED to the detected fluxes (as a grey shaded area in Figs. 4 and A.1), we only find the 70 μm flux to be higher for some of the objects classified as transitional disks candidates. Hence, we cannot conclude that we have detected a trend in the transitional disk population, such as has been observed in previous investigations (Ribas et al. 2013, Bustamante et al. 2015). We also find the case of 2MASS J16285694-2431096 (#18), where the 70 μm flux is not only lower than the median SED, but lower than the third quartile. Because of the lack of large excess in mid- to far-IR and the fact that this object presents a slightly smaller flux in the 70 μm band than the rest of transitional disks candidates, it needs further study to clarify its nature.

5. Conclusions

We have detected 49 objects in the central region of Ophiuchus in at least one PACS band and 19 in at least one SPIRE band. We obtained accurate photometric fluxes for the detected objects by means of aperture photometry.

Seven of the detected objects were classified as transitional disk candidates by the criterion in Ribas et al. (2013), and 11 more were classified according to the complementary criterion, generating a total sample of 18 transitional disk candidates. Some of these candidates were already imaged in previous works, and hence, confirmed. Six more objects are added to the final classification of candidates that have transitional features in their SEDs, rather than fulfilling any of the criterion. All of the transitional disk candidates are shown in Table 6 along with their classification criteria. This large difference between the identification methods can be due to the different nature or evolutionary stages of the disks, creating different geometries and leading to a large diversity of SEDs.

Several of the objects classified as transitional disks candidates have not been considered candidates before, but 2MASS J16281385-2456113 (#14) appears to be an attractive object for follow-up because of its prominent change of slope when compared to the rest of the sample, including previously imaged disks, such as IRS 48 (#9).

So far, *Herschel* data has proved to be very useful to improve the characterization of the outer regions of protoplanetary systems because of its long-wavelength coverage, unattainable until now, and its improved sensitivity and spatial resolution compared with previous IR missions. A study of the SED population of the disk sample detected with *Herschel* should give us more information on the true nature of these disks, but this study is outside the scope of this work.

Acknowledgements. We thank the referee for his/her constructive comments. This work has been possible thanks to the ESAC Space Science Faculty for funding with code ESAC-321, ESAC Traineeship program and of the Herschel Science Centre. We thank Gábor Marton for useful discussions on Herschel data evaluation. This work was partly supported by the Hungarian research grant OTKA 101393, and by the Momentum grant of the MTA CSFK Lendület Disk Research Group. PACS was developed by a consortium of institutes led by MPE (Germany), including UVIE (Austria); KUL, CSL, IMEC (Belgium); CEA, OAMP (France); MPIA (Germany); IFSI, OAP/AOT, OAA/CAISMI, LENS, SISSA (Italy); IAC (Spain). This development has been supported by the funding agencies BMVIT (Austria), ESA-PRODEX (Belgium), CEA/CNRS (France), DLR (Germany), ASI (Italy), and CICT/MCT (Spain). SPIRE was developed by a consortium of institutes led by Cardiff Univ. (UK), including Univ. Lethbridge (Canada); NAOC (China); CEA, LAM (France); IFSI, Univ. Padua (Italy); IAC (Spain); Stockholm Observatory (Sweden); Imperial College London, RAL, UCL-MSSL, UKATC, Univ. Sussex (UK); and Caltech, JPL, NHSC, Univ. Colorado (USA). This development has been supported by national funding agencies: CSA (Canada); NAOC (China); CEA, CNRS, CNRS (France); ASI (Italy); MCINN (Spain); SNSB (Sweden); STFC (UK); and NASA (USA). This study also makes use of the data products from the Two Micron All Sky Survey (2MASS), a joint project of the University of Massachusetts and IPAC/Caltech, funded by NASA and the National Science Foundation; data products from the Wide-field Infrared Survey Explorer (WISE), a joint project of the University of California, Los Angeles, and the Jet Propulsion Laboratory (JPL)/California Institute of Technology (Caltech); data products from DENIS, a project partly funded by the SCIENCE and the HCM plans of the European Commission under grants CT920791 and CT940627; the NASA Infrared Processing and Analysis Center (IPAC) Science Archive; and the SIMBAD database.

References

- ALMA Partnership, Brogan, C. L., Perez, L. M., et al. 2015, *ApJ*, **808**, L3
 Alves de Oliveira, C., Moraux, E., Bouvier, J., et al. 2010, *A&A*, **515**, A75
 Alves de Oliveira, C., Ábrahám, P., Marton, G., et al. 2013, *A&A*, **559**, A126
 André, P., Men'shchikov, A., Bontemps, S., et al. 2010, *A&A*, **518**, L102

- Andrews, S. M., Wilner, D. J., Hughes, A. M., Qi, C., & Dullemond, C. P. 2009, *ApJ*, **700**, 1502
- Andrews, S. M., Wilner, D. J., Espaillat, C., et al. 2011, *ApJ*, **732**, 42
- Balog, Z., Müller, T., Nielbock, M., et al. 2014, *Exp. Astron.*, **37**, 129
- Bustamante, I., Merín, B., Ribas, Á., et al. 2015, *A&A*, **578**, A23
- Cieza, L. A., Schreiber, M. R., Romero, G. A., et al. 2010, *ApJ*, **712**, 925
- Clarke, C. J., Gendrin, A., & Sotomayor, M. 2001, *MNRAS*, **328**, 485
- Dullemond, C. P., & Dominik, C. 2005, *A&A*, **434**, 971
- Erickson, K. L., Wilking, B. A., Meyer, M. R., Robinson, J. G., & Stephenson, L. N. 2011, *AJ*, **142**, 140
- Espaillat, C., Muzerolle, J., Najita, J., et al. 2014, in *Protostars and Planets VI*, eds. R. S. Klessen, C. P. Dullemond & T. Henning (Tucson: University of Arizona Press), 497
- Evans, II, N. J., Dunham, M. M., Jørgensen, J. K., et al. 2009, *ApJS*, **181**, 321
- Fazio, G. G., Hora, J. L., Allen, L. E., et al. 2004, *ApJS*, **154**, 10
- Griffin, M. J., Abergel, A., Abreu, A., et al. 2010, *A&A*, **518**, L3
- Hauschildt, P. H., Allard, F., & Baron, E. 1999, *ApJ*, **512**, 377
- Ho, P. T. P., Moran, J. M., & Lo, K. Y. 2004, *ApJ*, **616**, L1
- Kim, K. H., Watson, D. M., Manoj, P., et al. 2013, *ApJ*, **769**, 149
- Lada, C. J. 1987, in *Star Forming Regions*, eds. M. Peimbert, & J. Jugaku, *IAU Symp.*, **115**, 1
- Marsh, K. A., & Mahoney, M. J. 1992, *ApJ*, **395**, L115
- Natta, A., Testi, L., Comerón, F., et al. 2002, *A&A*, **393**, 597
- Ott, S. 2010, in *Astronomical Data Analysis Software and Systems XIX*, eds. Y. Mizumoto, K.-I. Morita, & M. Ohishi, *ASP Conf. Ser.*, **434**, 139
- Pilbratt, G. L., Riedinger, J. R., Passvogel, T., et al. 2010, *A&A*, **518**, L1
- Poglitsch, A., Waelkens, C., Geis, N., et al. 2010, *A&A*, **518**, L2
- Ribas, Á., Merín, B., Bouy, H., et al. 2013, *A&A*, **552**, A115
- Ribas, Á., Merín, B., Bouy, H., & Maud, L. T. 2014, *A&A*, **561**, A54
- Rieke, G. H., Young, E. T., Engelbracht, C. W., et al. 2004, *ApJS*, **154**, 25
- Roussel, H. 2013, *PASP*, **125**, 1126
- Skrutskie, M. F., Cutri, R. M., Stiening, R., et al. 2006, *AJ*, **131**, 1163
- Strom, K. M., Strom, S. E., Edwards, S., Cabrit, S., & Skrutskie, M. F. 1989, *AJ*, **97**, 1451
- Traficante, A., Fuller, G. A., Pineda, J. E., & Pezzuto, S. 2015, *A&A*, **574**, A119
- van der Marel, N., van Dishoeck, E. F., Bruderer, S., et al. 2013, *Science*, **340**, 1199
- Weingartner, J. C., & Draine, B. T. 2001, *ApJ*, **548**, 296
- Werner, M. W., Roellig, T. L., Low, F. J., et al. 2004, *ApJS*, **154**, 1
- Wilking, B. A., Meyer, M. R., Robinson, J. G., & Greene, T. P. 2005, *AJ*, **130**, 1733
- Wilking, B. A., Gagné, M., & Allen, L. E. 2008, *Star Formation in the ρ Ophiuchi Molecular Cloud*, ed. B. Reipurth, 351
- Wooten, A., & Thompson, A. R. 2009, *IEEE Proc.*, **97**, 1463
- Wright, E. L., Eisenhardt, P. R. M., Mainzer, A. K., et al. 2010, *AJ*, **140**, 1868

Appendix A

Table A.1. Estimation of contaminating flux contained in the aperture for each band, according to the median SED extrapolation.

ID	Name	F_{70} [Jy] (6'')	F_{160} [Jy] (12'')	F_{250} [Jy] (22'')	F_{350} [Jy] (30'')	F_{500} [Jy] (42'')
–	2MASS J16245974-2456008	–	2.66×10^{-5}	2.94×10^{-4}	9.00×10^{-5}	2.77×10^{-5}
–	V*V852Oph	–	1.44×10^{-4}	1.48×10^{-4}	4.96×10^{-5}	6.11×10^{-5}
–	2MASS J16253673-2415424	–	–	4.47×10^{-5}	1.23×10^{-5}	3.17×10^{-5}
–	2MASS J16253958-2426349	–	–	9.65×10^{-4}	4.60×10^{-4}	3.15×10^{-4}
–	V*V2058Oph	3.34×10^{-2}	1.12×10^{-2}	5.29×10^{-3}	1.46×10^{-3}	3.97×10^{-4}
1	Haro1-6	5.85×10^{-2}	1.55×10^{-2}	7.64×10^{-3}	2.15×10^{-3}	6.09×10^{-4}
–	2MASS J16261684-2422231	–	–	1.01×10^{-5}	1.15×10^{-5}	1.48×10^{-5}
2	DoAr 24	4.51×10^{-3}	9.56×10^{-4}	4.50×10^{-4}	1.32×10^{-4}	4.03×10^{-5}
–	2MASS J16262189-2444397	5.90×10^{-3}	1.27×10^{-3}	5.98×10^{-4}	2.07×10^{-4}	7.49×10^{-5}
–	DoAr 24E	–	–	–	1.41×10^{-4}	3.85×10^{-5}
3	DoAr 25	–	1.23×10^{-4}	8.50×10^{-5}	3.09×10^{-5}	2.91×10^{-5}
–	GSS32	–	3.76×10^{-2}	1.77×10^{-2}	5.67×10^{-3}	1.61×10^{-3}
–	2MASS J16262407-2416134	–	–	–	6.89×10^{-5}	2.34×10^{-5}
–	2MASS J16263297-2400168	–	–	1.32×10^{-5}	1.86×10^{-5}	1.06×10^{-5}
–	2MASS J16263682-2415518	–	1.39×10^{-3}	6.68×10^{-4}	2.07×10^{-4}	6.18×10^{-5}
–	[GY92]93	–	–	1.36×10^{-5}	1.12×10^{-5}	1.59×10^{-5}
–	2MASS J16264285-2420299	–	–	–	2.40×10^{-4}	1.60×10^{-4}
–	2MASS J16264643-2412000	–	3.32×10^{-4}	1.65×10^{-4}	4.94×10^{-5}	1.69×10^{-5}
4	WSB 40	2.09×10^{-2}	4.43×10^{-3}	2.09×10^{-3}	5.86×10^{-4}	1.65×10^{-4}
–	WL18	–	–	4.75×10^{-6}	5.77×10^{-6}	5.37×10^{-6}
–	2MASS J16265677-2413515	1.59×10^{-3}	3.36×10^{-4}	1.58×10^{-4}	5.43×10^{-5}	2.13×10^{-5}
5	SR 24S	3.92×10^{-1}	8.29×10^{-2}	3.91×10^{-2}	1.08×10^{-2}	2.95×10^{-3}
6	2MASS J16270659-2441488	–	2.93×10^{-5}	2.17×10^{-5}	1.28×10^{-5}	1.23×10^{-5}
–	2MASS J16270907-2412007	–	–	2.68×10^{-5}	1.86×10^{-5}	8.30×10^{-6}
7	WSB 46	–	4.76×10^{-5}	2.46×10^{-5}	1.65×10^{-5}	8.68×10^{-6}
–	[WMR2005]4 – 10	5.65×10^{-4}	1.20×10^{-4}	7.63×10^{-5}	5.06×10^{-5}	2.52×10^{-5}
–	2MASS J16271836-2454537	–	–	7.27×10^{-6}	3.56×10^{-6}	3.40×10^{-6}
–	WSB 49	–	–	1.78×10^{-5}	6.58×10^{-6}	4.63×10^{-6}
–	2MASS J16272658-2425543	–	–	4.66×10^{-5}	3.35×10^{-5}	1.53×10^{-5}
8	2MASS J16273084-2424560	–	–	1.72×10^{-5}	1.30×10^{-5}	5.67×10^{-6}
–	2MASS J16273311-2441152	–	–	–	3.20×10^{-6}	5.10×10^{-6}
9	IRS 48	–	–	–	1.20×10^{-4}	4.42×10^{-5}
10	IRS 49	–	–	6.90×10^{-6}	7.97×10^{-6}	6.54×10^{-6}
–	2MASS J16273832-2357324	–	–	7.64×10^{-4}	2.19×10^{-4}	6.12×10^{-5}
11	2MASS J16273863-2438391	6.27×10^{-3}	1.62×10^{-3}	5.83×10^{-5}	3.59×10^{-5}	1.51×10^{-3}
–	2MASS J16273901-2358187	–	–	1.51×10^{-4}	5.32×10^{-5}	2.76×10^{-5}
12	WSB52	1.45×10^{-3}	3.08×10^{-4}	1.56×10^{-4}	1.61×10^{-5}	1.25×10^{-4}
13	SR 9	–	–	1.70×10^{-5}	1.59×10^{-5}	1.02×10^{-5}
–	2MASS J16274270-2438506	–	–	5.44×10^{-6}	4.54×10^{-6}	6.99×10^{-6}
–	V*V2059Oph	–	–	–	8.23×10^{-6}	7.76×10^{-6}
–	2MASS J16280256-2355035	–	–	6.06×10^{-8}	5.08×10^{-6}	5.56×10^{-6}
–	2MASS J16281379-2432494	–	–	5.82×10^{-6}	4.25×10^{-6}	7.35×10^{-6}
14	2MASS J16281385-2456113	–	–	1.02×10^{-4}	2.94×10^{-5}	1.63×10^{-5}
15	WSB 60	–	2.13×10^{-4}	1.10×10^{-5}	1.00×10^{-5}	6.31×10^{-6}
–	2MASS J16281673-2405142	–	–	–	7.31×10^{-6}	8.51×10^{-6}
16	SR 20W	–	–	4.36×10^{-6}	5.78×10^{-6}	6.48×10^{-6}
17	SR 13	–	–	–	4.01×10^{-5}	1.51×10^{-5}
18	2MASS J16285694-2431096	–	–	4.67×10^{-6}	1.83×10^{-5}	7.36×10^{-6}
–	2MASS J16294427-2441218	–	–	1.75×10^{-5}	1.35×10^{-5}	7.44×10^{-6}

Notes. The contaminating flux has been calculated extrapolating the MIPS-24 flux obtained from the *Spitzer* c2d catalogue (Evans et al. 2009) to the median SED for each source contained in the different apertures for each band.

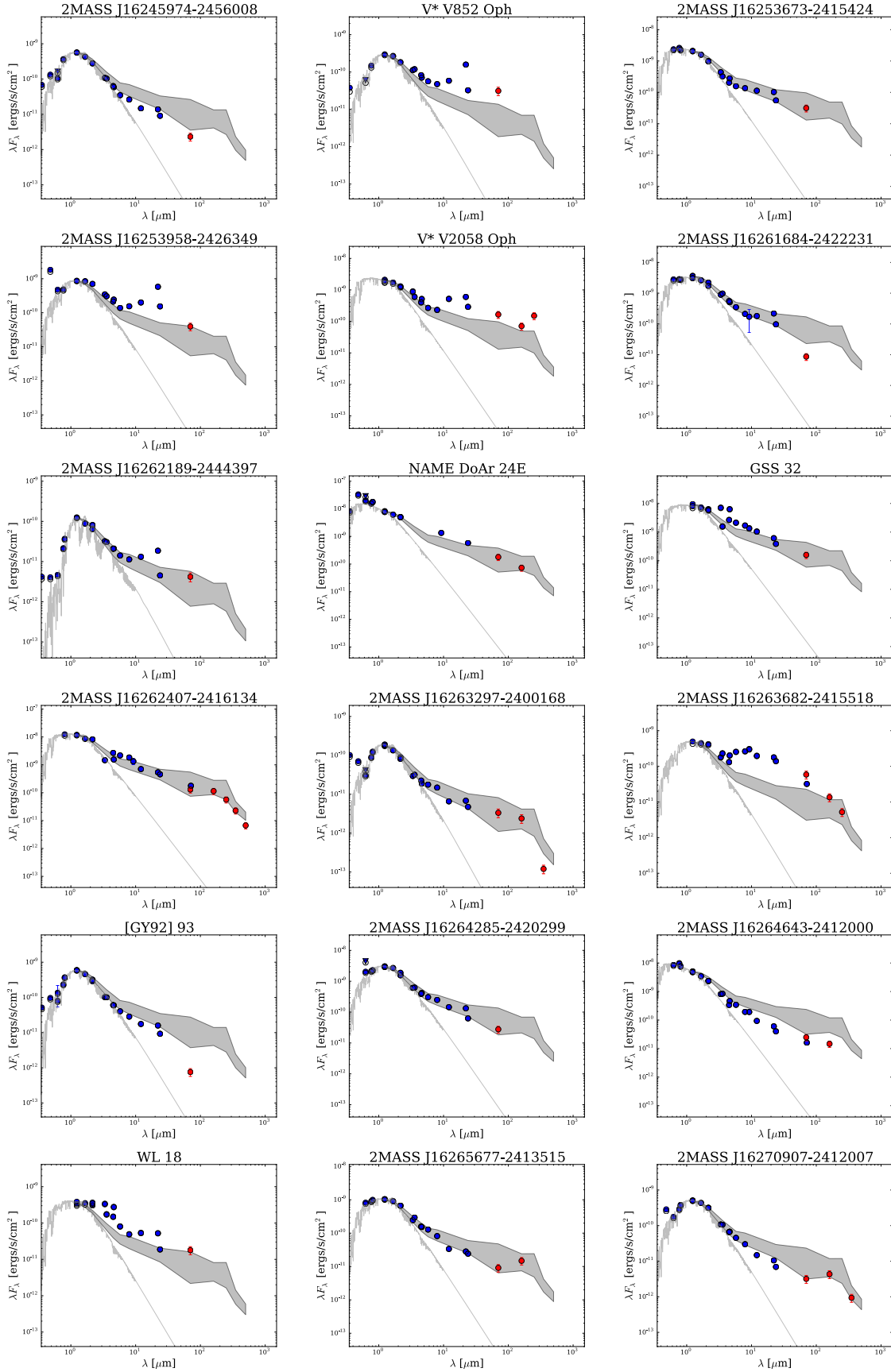


Fig. A.1. Spectral energy distribution (SED) of the sources detected in at least one band by *Herschel* and classified as non-transitional. Blue dots show data acquired from the literature, red dots are photometric fluxes obtained from *Herschel* data. Grey dashed line is the photosphere model according to the spectral type, and the grey shaded area is the filled area between the first and third quartile of all the disk fluxes. Observed fluxes are shown with empty circles and A_v values used are in Table 3.

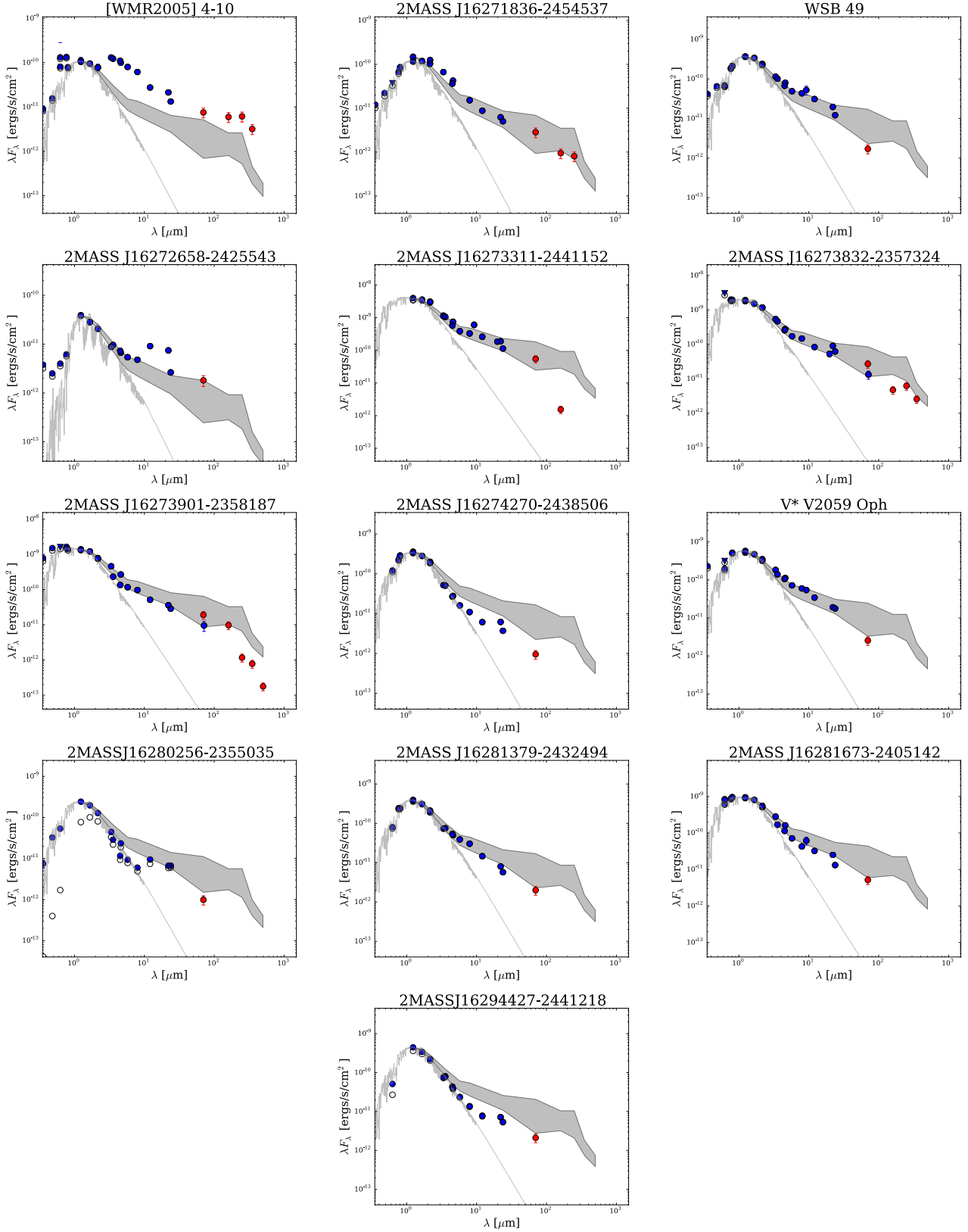


Fig. A.1. continued.

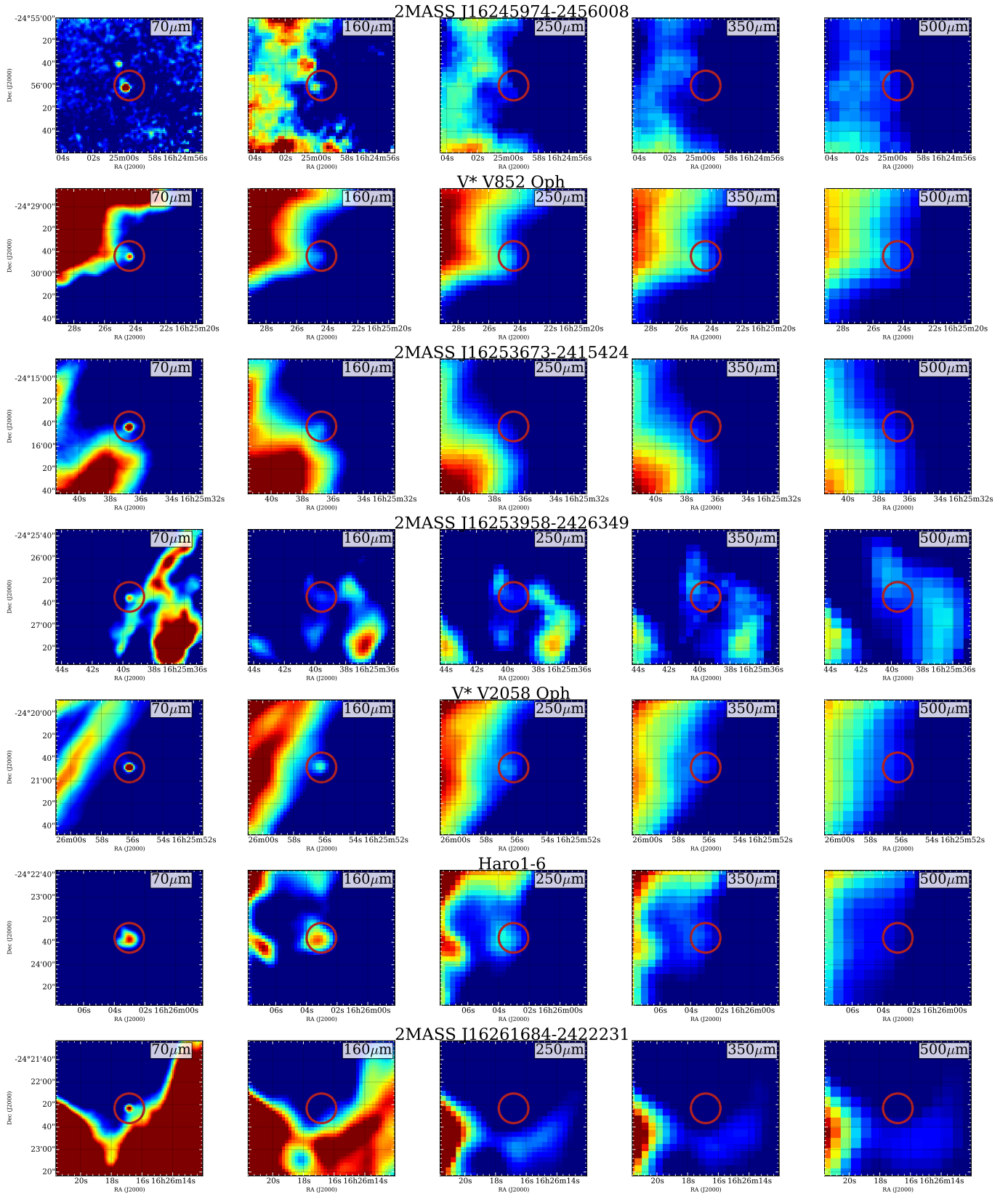


Fig. A.2. Thumbnail images of each of the 46 sources with at least one point source detected by *Herschel*. All images are $60'' \times 60''$ with north up and east to the left. The cut levels are set to the rms of the background for the minimum and three times that value for the maximum with linear stretch.

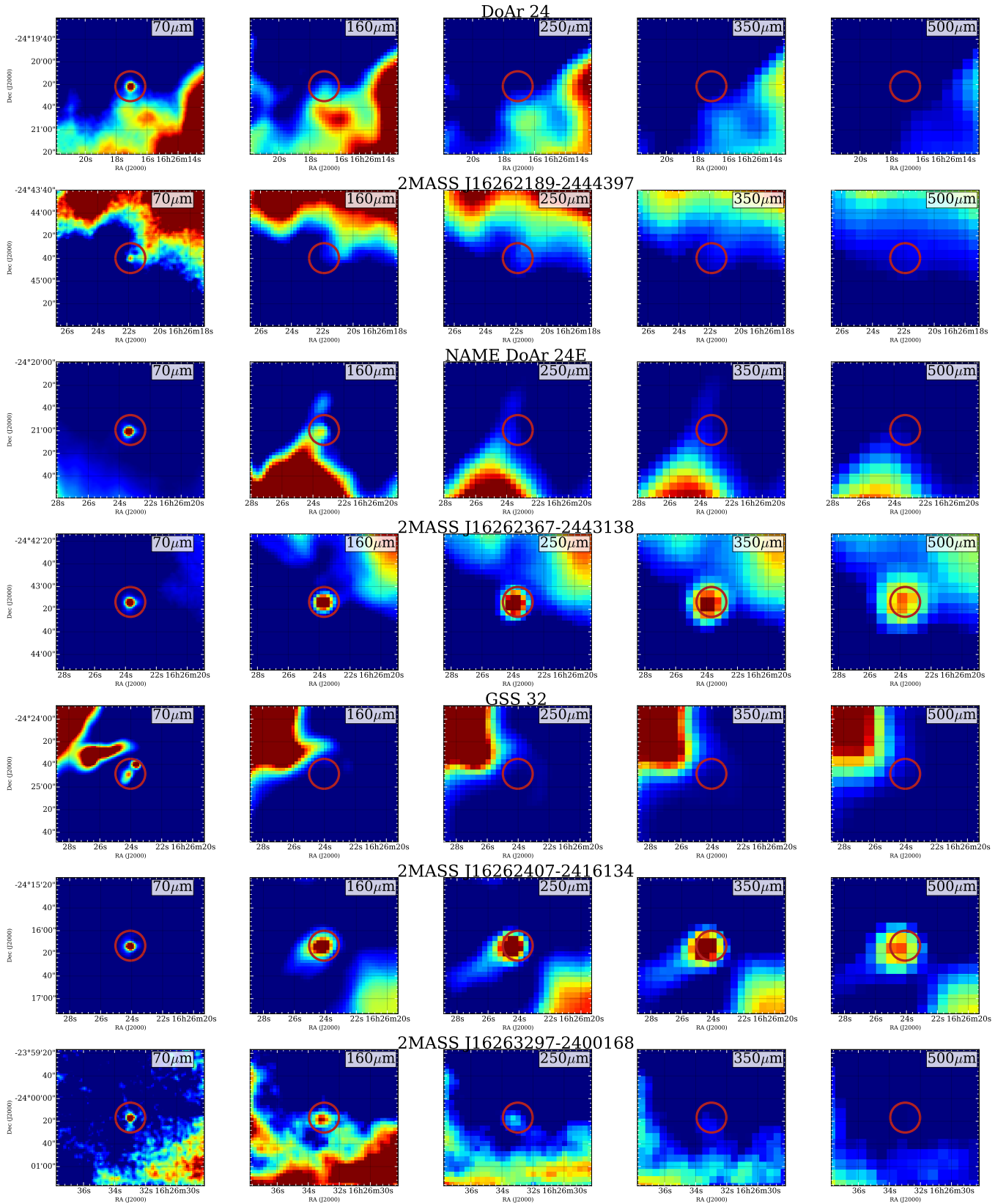
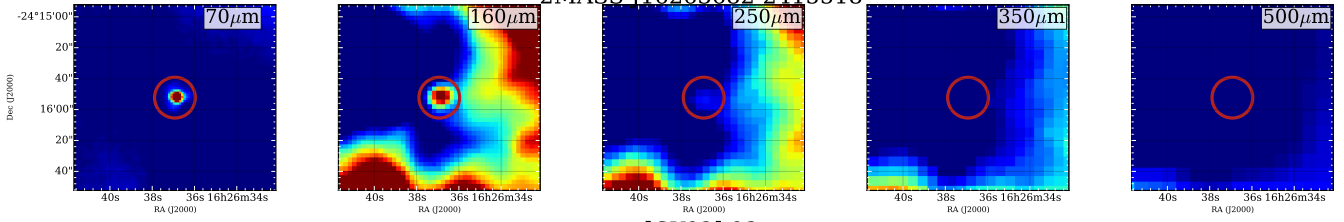
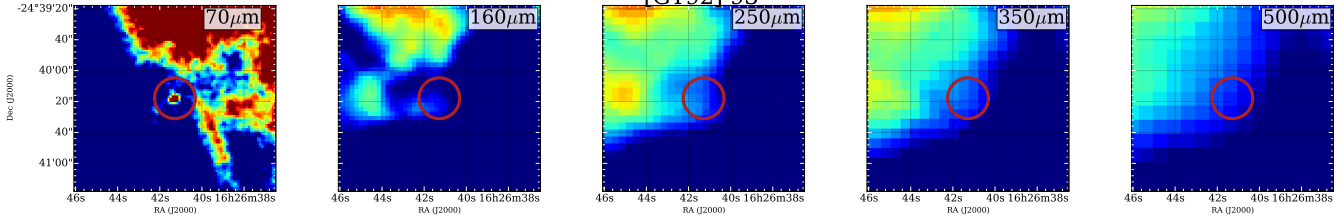


Fig. A.2. continued.

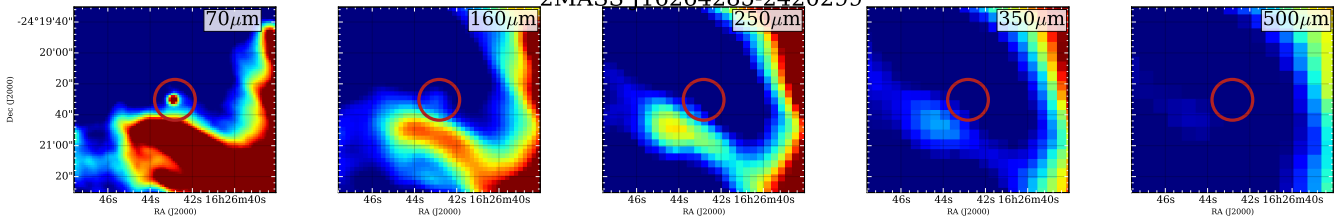
2MASS J16263682-2415518



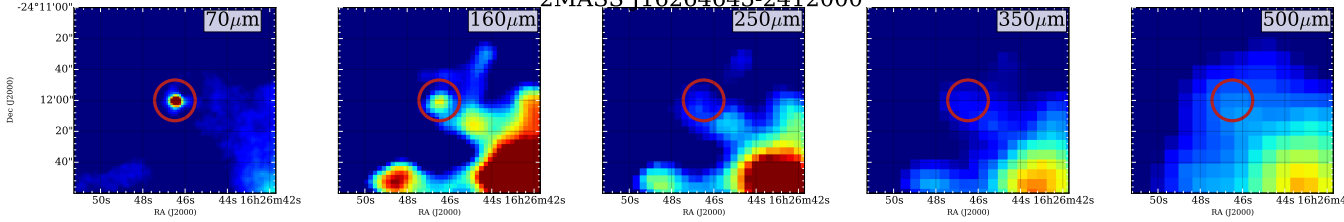
[GY92] 93



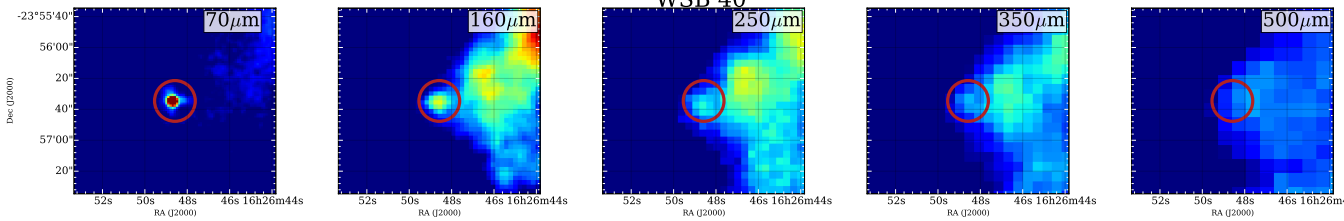
2MASS J16264285-2420299



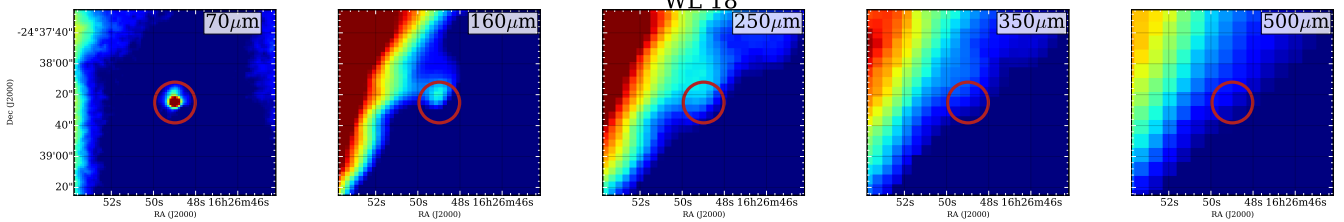
2MASS J16264643-2412000



WSB 40



WL 18



2MASS J16265677-2413515

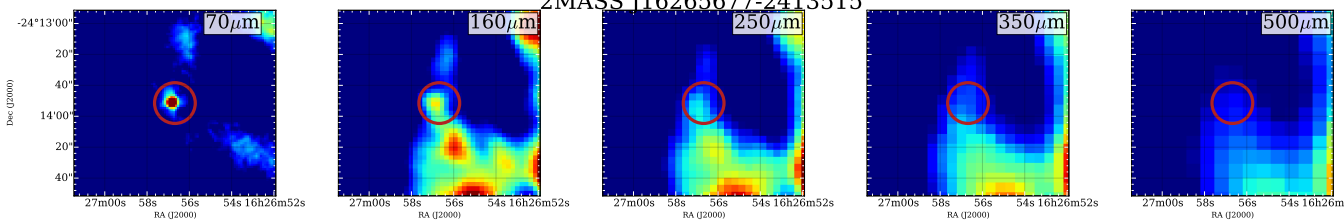


Fig. A.2. continued.

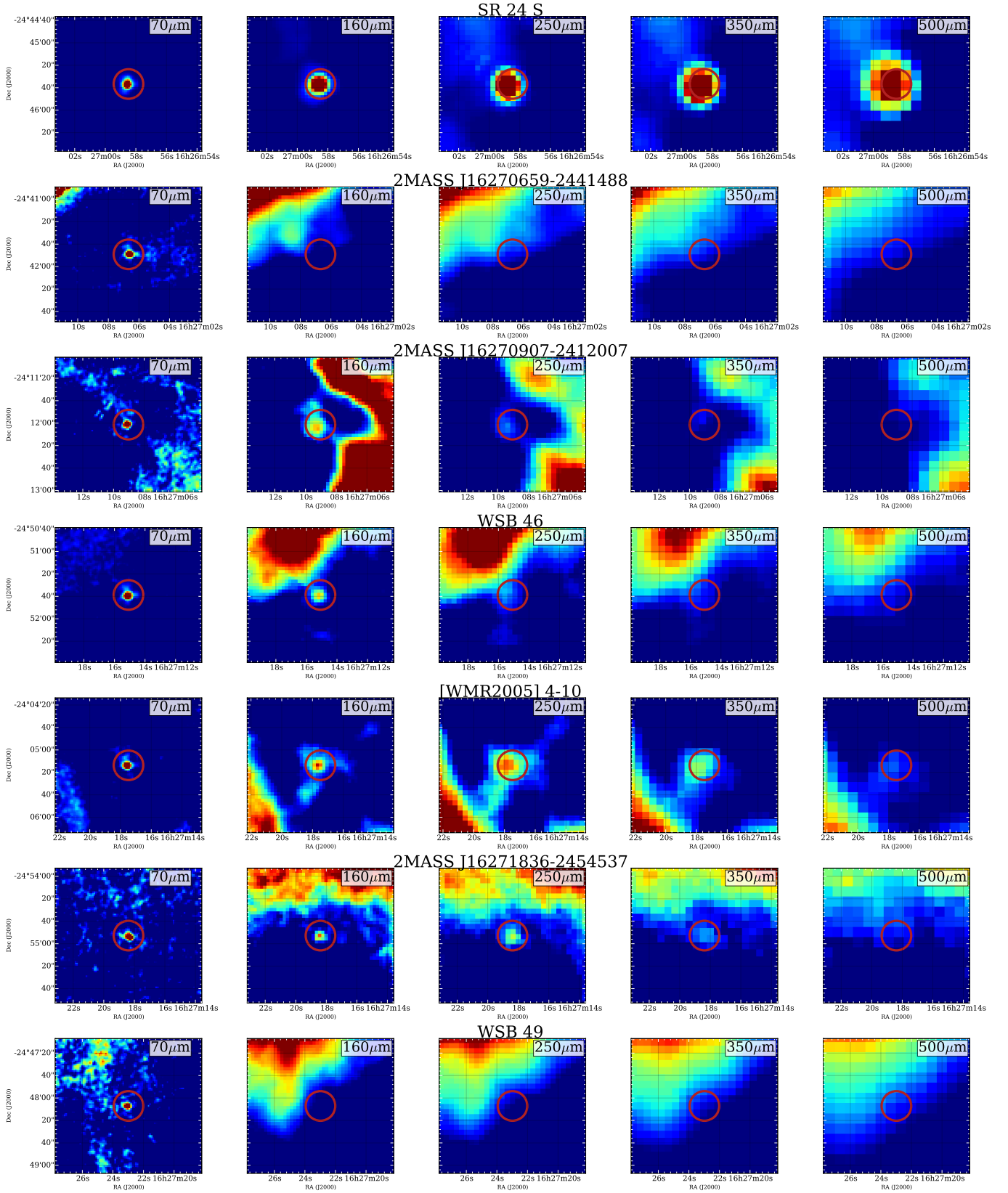


Fig. A.2. continued.

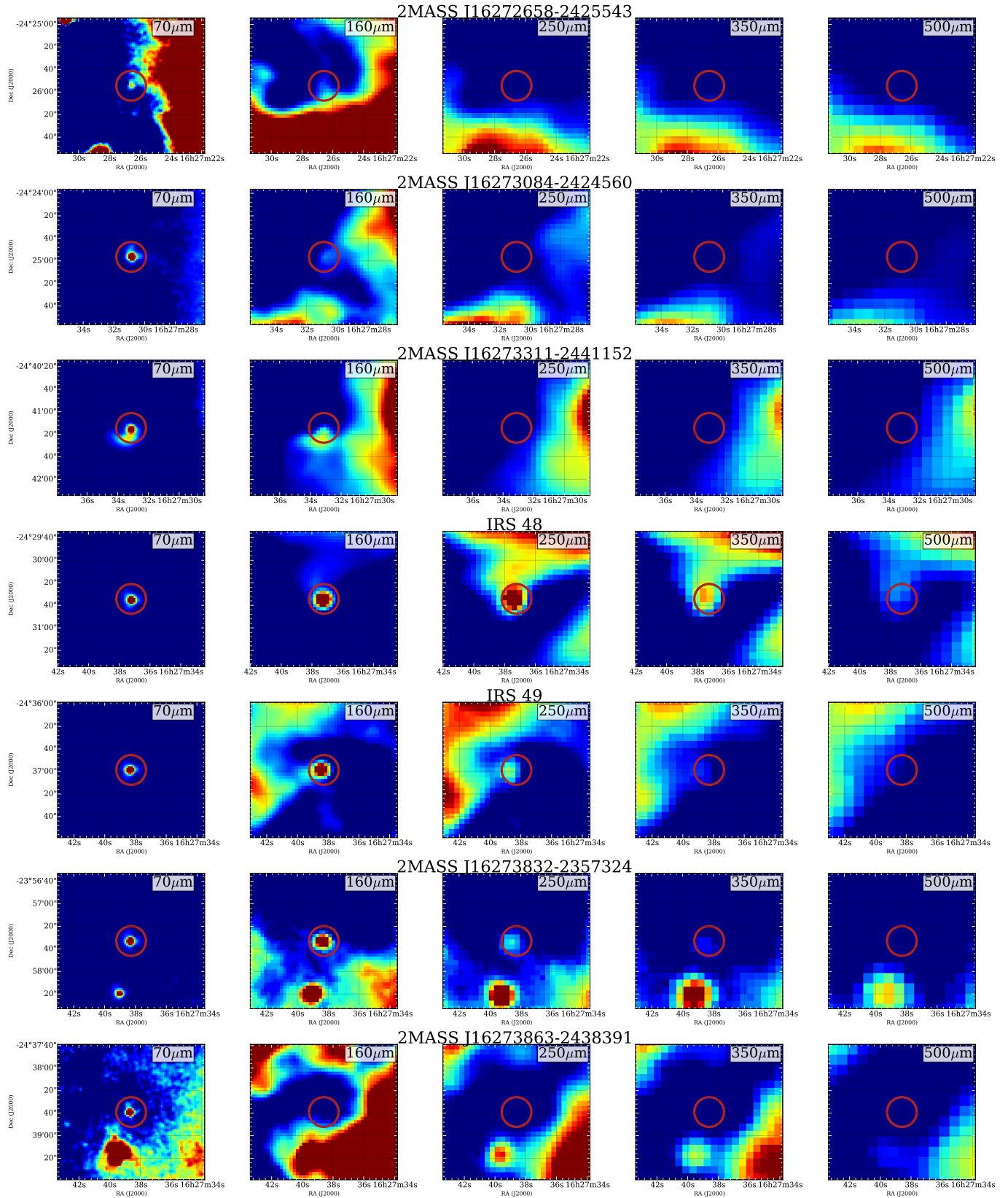


Fig. A.2. continued.

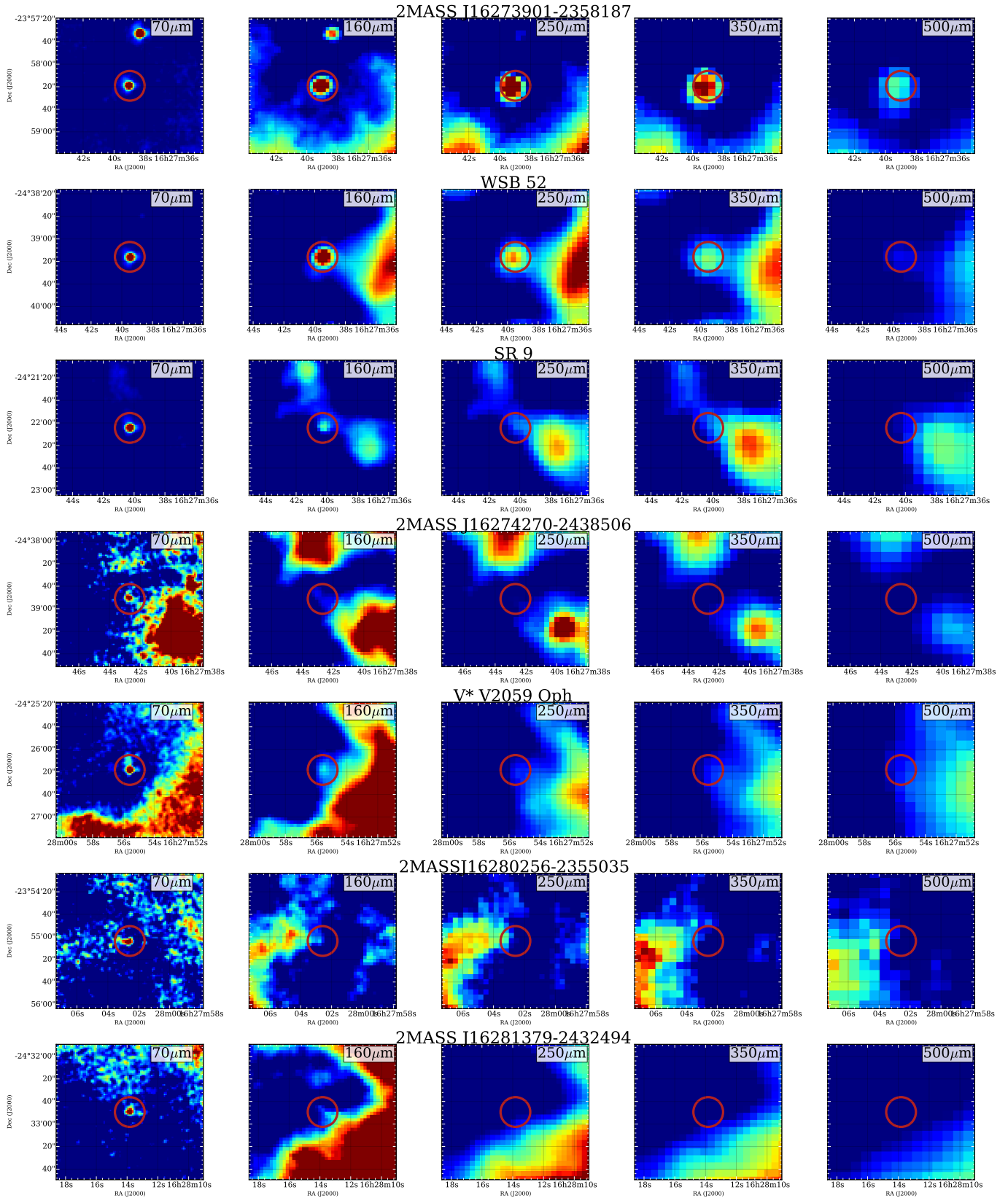
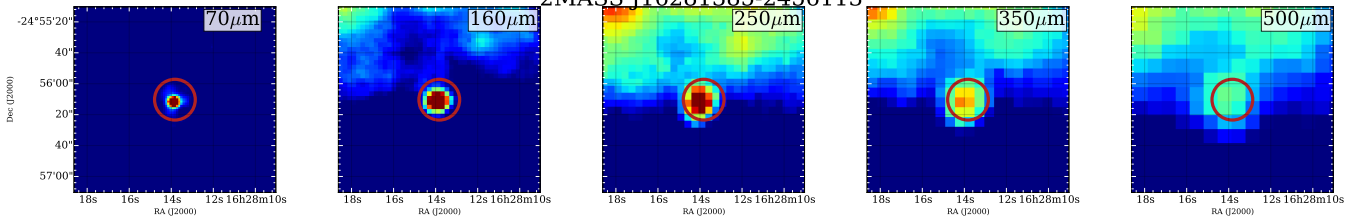
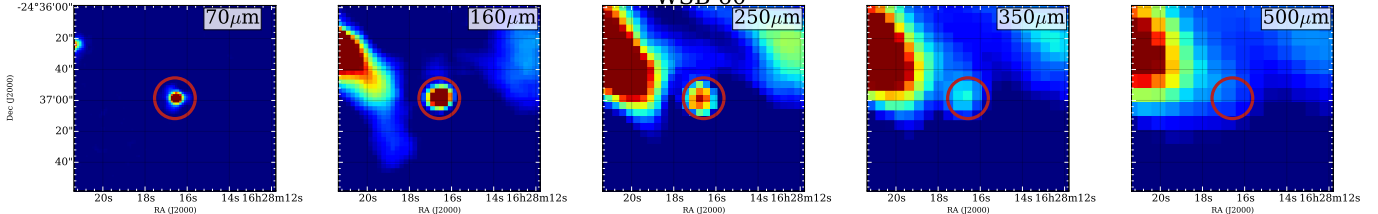


Fig. A.2. continued.

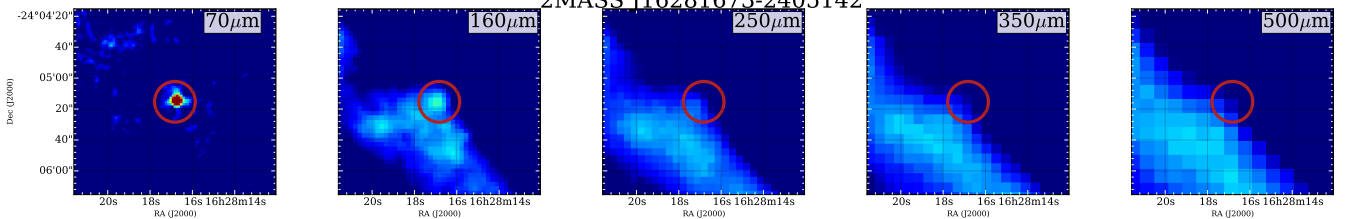
2MASS J16281385-2456113



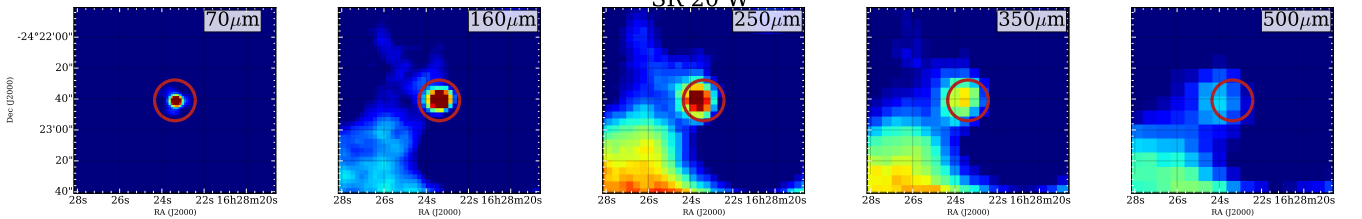
WSB 60



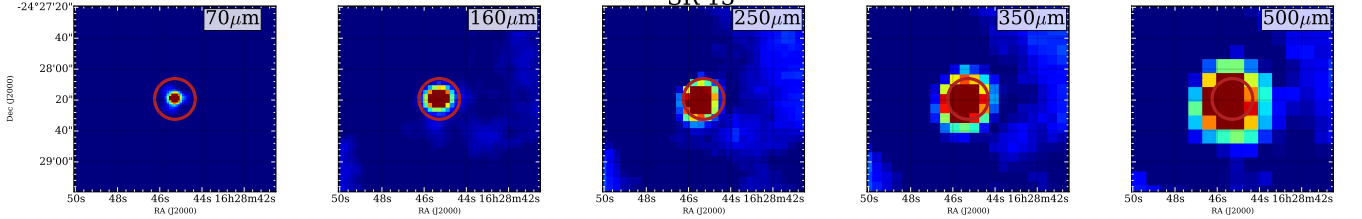
2MASS J16281673-2405142



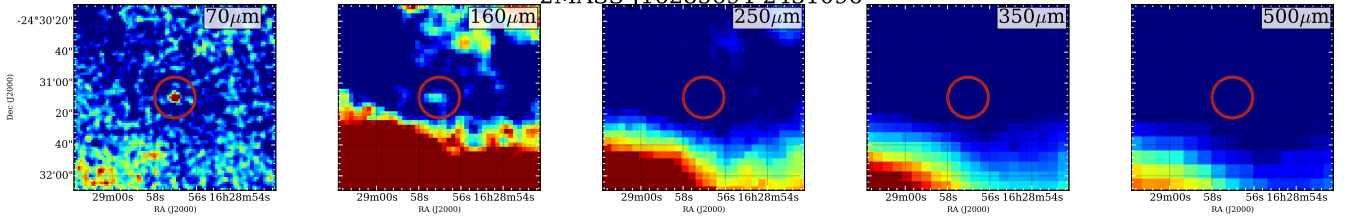
SR 20 W



SR 13



2MASS J16285694-2431096



2MASS J16294427-2441218

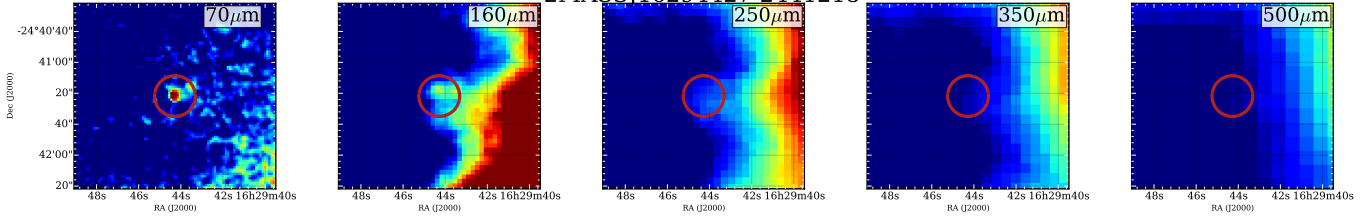


Fig. A.2. continued.