



Measuring cluster masses with CMB lensing: a statistical approach

Jean-Baptiste Melin, James G. Bartlett

► To cite this version:

Jean-Baptiste Melin, James G. Bartlett. Measuring cluster masses with CMB lensing: a statistical approach. *Astronomy and Astrophysics - A&A*, 2015, 578, pp.A21. 10.1051/0004-6361/201424720 . cea-01300573

HAL Id: cea-01300573

<https://cea.hal.science/cea-01300573>

Submitted on 11 Apr 2016

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

Measuring cluster masses with CMB lensing: a statistical approach

Jean-Baptiste Melin¹ and James G. Bartlett^{2,3}

¹ DSM/Irfu/SPP, CEA-Saclay, 91191 Gif-sur-Yvette Cedex, France
e-mail: jean-baptiste.melin@cea.fr

² APC, AstroParticule et Cosmologie, Université Paris Diderot, CNRS/IN2P3, CEA/Irfu, Observatoire de Paris, Sorbonne Paris Cité,
10 rue Alice Domon et Léonie Duquet, 75205 Paris Cedex 13, France
e-mail: bartlett@apc.univ-paris7.fr

³ Jet Propulsion Laboratory, California Institute of Technology, 4800 Oak Grove Drive, Pasadena, California, USA

Received 31 July 2014 / Accepted 10 April 2015

ABSTRACT

We present a method for measuring the masses of galaxy clusters using the imprint of their gravitational lensing signal on the cosmic microwave background (CMB) temperature anisotropies. The method first reconstructs the projected gravitational potential with a quadratic estimator and then applies a matched filter to extract cluster mass. The approach is well-suited for statistical analyses that bin clusters according to other mass proxies. We find that current experiments, such as *Planck*, the South Pole Telescope and the Atacama Cosmology Telescope, can practically implement such a statistical methodology, and that future experiments will reach sensitivities sufficient for individual measurements of massive systems. As illustration, we use simulations of *Planck* observations to demonstrate that it is possible to constrain the mass scale of a set of 62 massive clusters with prior information from X-ray observations, similar to the published *Planck* ESZ-XMM sample. We examine the effect of the thermal (tSZ) and kinetic (kSZ) Sunyaev-Zeldovich (SZ) signals, finding that the impact of the kSZ remains small in this context. The stronger tSZ signal, however, must be actively removed from the CMB maps by component separation techniques prior to reconstruction of the gravitational potential. Our study of two such methods highlights the importance of broad frequency coverage for this purpose. A companion paper presents application to the *Planck* data on the ESZ-XMM sample.

Key words. large-scale structure of Universe – galaxies: clusters: general – cosmic background radiation – methods: data analysis – gravitational lensing: weak – methods: statistical

1. Introduction

The most important property of galaxy clusters for cosmological studies is their mass; it is also the most difficult to measure, because it is not directly observable. Accurate mass measurements are needed, in particular, to calibrate scaling laws relating mass to observable cluster properties, such as richness (Yee & Ellingson 2003; Gladders et al. 2007; Rozo et al. 2009), X-ray properties (Arnaud et al. 2005, 2007; Stanek et al. 2006; Vikhlinin et al. 2006; Pratt et al. 2009; Mantz et al. 2010; Rozo et al. 2014b) or Sunyaev-Zeldovich (SZ) signal strength (Marrone et al. 2012; Planck Collaboration Int. III 2013; Bocquet et al. 2015; Rozo et al. 2014a; von der Linden et al. 2014; Hoekstra et al. 2015). The uncertainty in the mass calibration of these relations now limits the constraining power of cluster counts as a cosmological probe (Roza et al. 2013; Planck Collaboration XX 2014; Hasselfield et al. 2013; Reichardt et al. 2013; Roza et al. 2010; Mantz et al. 2010, 2015; Vikhlinin et al. 2009; Böhringer et al. 2014; Planck Collaboration 2015).

Cluster masses can be determined dynamically, by application of the virial theorem to the velocity distribution of member galaxies, from X-ray studies of the intra-cluster medium, assuming hydrostatic equilibrium for the gas, and via the effects of gravitational lensing that distort the shapes of background galaxies. Each approach presents its own advantages while suffering from specific systematic biases (Allen et al. 2011).

In this work, we discuss lensing of the cosmic microwave background (CMB) anisotropies as a promising new technique for measuring cluster masses, presenting a methodology for practical application with the specific aim of calibrating cluster scaling relations. We evaluate the potential of current and future experiments to employ the methodology and show how to account for astrophysical biases from other cluster signals. As illustration, we apply the technique to *Planck* simulations of massive clusters and demonstrate that it is possible to recover an unbiased estimate of the mass scale of the *Planck* XMM-Early Release SZ catalogue (ESZ-XMM, Planck Collaboration VIII 2011; Planck Collaboration XI 2011). A companion paper presents results obtained with the *Planck* dataset for the ESZ-XMM sample.

Study of CMB lensing (for a review see Lewis & Challinor 2006) is a rapidly growing field driven by the current generation of sensitive, high resolution CMB experiments. Recent measurements of lensing in the CMB temperature power spectrum have been given by the Atacama Cosmology Telescope (ACT, Das et al. 2014), the South Pole Telescope (SPT, Story et al. 2013) and the *Planck* mission (Planck Collaboration XVI 2014). Direct reconstruction of the matter power spectrum using higher order statistics, designed to capture lensing's specific non-Gaussian mode coupling signature (Hu & Okamoto 2002), have also been reported by ACT (Das et al. 2011) and SPT (van Engelen et al. 2012), as well as *Planck*, which in addition produced an all-sky map of the projected matter distribution (Planck Collaboration XVII 2014).

Lensing of the CMB by galaxy clusters was first discussed at length by Zaldarriaga & Seljak (1999) when developing methods for reconstruction of the gravitational lensing potential. Seljak & Zaldarriaga (2000) considered instead the characteristic perturbation to the unlensed CMB temperature field, approximated locally as a pure gradient, induced by cluster lensing, an idea further studied by Holder & Kosowsky (2004) and Vale et al. (2004). In an approach analogous to the first authors, Maturi et al. (2005) built a filter nonlinear in CMB temperature to reconstruct the lensing convergence field around clusters as a means of studying their density profiles.

In this paper, we develop a cluster mass extraction method based on a matched filter for the projected gravitational potential. Our approach is similar in spirit to the work of Zaldarriaga & Seljak (1999) and Maturi et al. (2005) in that we first reconstruct the lensing field around a cluster, rather than working with the lensing perturbation in the CMB temperature itself. In the present work, however, we focus on determining total cluster mass even in relatively low signal-to-noise regimes. Once we have the map of a cluster potential, obtained using the quadratic estimator of Hu & Okamoto (2002), we then apply a matched filter designed to optimally extract the cluster's mass, assuming a density profile. This allows us to obtain measurements for individual clusters, even when noise dominated, and use them in statistical analyses; for example, finding the mean mass of clusters by binning according to SZ signal strength or, in other words, the mass-SZ scaling relation.

In practice, we must confront a number of possible systematic biases. We consider the effects of and ways to mitigate astrophysical signals contaminating the CMB map required as input for reconstruction of the gravitational potential. The most difficult in this context are signals generated by the cluster itself, such as the thermal SZ (tSZ, Sunyaev & Zeldovich 1972) and the kinetic SZ (kSZ, Sunyaev & Zeldovich 1980) effects. Component separation is therefore a crucial step prior to reconstruction of the potential, and our study will demonstrate the importance of multi-frequency observations in this context. The kSZ effect, having the same spectral signature as the CMB, requires separate treatment; fortunately, we will see that the lensing potential reconstruction significantly reduces its impact by averaging it with uncorrelated CMB anisotropies.

We organize the paper as follows. We begin in Sect. 2 by establishing our data model and discussing the reconstruction of the lensing potential and application of the matched filter; the presentation focuses at this point on the ideal case where we have a clean CMB map of known noise properties. This allows a preliminary evaluation of the potential of current and future CMB experiments to measure cluster masses. Section 3 focuses on the tSZ and kSZ signals. We employ two techniques to remove the tSZ signal and produce clean CMB maps from a set of individual frequency maps, and we evaluate the impact of the kSZ signal. We simulate *Planck* observations of a set of massive clusters in Sect. 4 to illustrate the method, showing that it is possible (Sect. 5) to recover an unbiased estimate of the cluster mass scale for a sample similar to the *Planck* ESZ-XMM catalogue (Planck Collaboration XI 2011). We finish with a discussion (Sect. 6) and our conclusions (Sect. 7). Throughout, we adopt a flat Λ CDM cosmological model with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $\Omega_M = 1 - \Omega_\Lambda = 0.3$.

2. Mass estimation

We first define our data model in the general context, and then present the matched filter mass estimation by focussing

on the ideal case where we have a clean map of lensed CMB anisotropies with only instrumental noise.

2.1. Cluster signals

Consider a patch of sky centered on a galaxy cluster of mass M_{500} at redshift z . We refer to cluster mass inside the radius R_{500} , interior to which the mean mass density is 500 times the critical density at the cluster's redshift, $\rho_c(z)$; i.e., $M_{500} = 500(4\pi/3)\rho_c(z)R_{500}^3$. The hot, gaseous intra-cluster medium (ICM) generates both tSZ and kSZ effects, and the cluster's (projected) gravitational potential, ϕ , lenses the CMB anisotropies by bending light rays and displacing the apparent line-of-sight.

We model the mass distribution with a NFW (Navarro et al. 1996) profile,

$$\rho(r) = \frac{\rho_0}{(r/r_s)(1 + r/r_s)^2}, \quad (1)$$

described by its central density, ρ_0 , and physical scale r_s . The latter can be related to R_{500} using the concentration parameter, c_{500} , as $r_s = R_{500}/c_{500}$. Unless otherwise stated, we take $c_{500} = 3$ in this work; in reality, it is expected to depend weakly on cluster mass and redshift, i.e., $c_{500}(M_{500}, z)$ (Bullock et al. 2001; Neto et al. 2007; Muñoz-Cuartas et al. 2011). Integrating along the line-of-sight yields the projected surface mass density at angular position x from the center,

$$\Sigma(x) = \phi_0 \Sigma_m(x/\theta_m), \quad (2)$$

where $x = |x|$ and Σ_m is a template characterized by the angular scale $\theta_m = r_s/D_{\text{clus}}$, with D_{clus} the angular diameter distance to the cluster. The integral along the line-of-sight is performed out to $r = 5R_{500}$. The normalization, ϕ_0 , is given below.

For the pressure of the ICM, we employ the modified NFW profile of Nagai et al. (2007) with parameters given by the X-ray observations of Arnaud et al. (2010), the so-called universal pressure profile in the non self-similar case (Eq. (12) of that paper). Integrating the pressure profile along the line-of-sight gives the tSZ angular template, $T_t(x/\theta_t)$, characterized by its scale radius, θ_t , and which is normalized by the cluster's central Compton y value, y_0 , to obtain the complete tSZ profile. The kSZ signal is proportional to the optical depth through the cluster with profile $T_k(x/\theta_k)$ characterized by the scale radius θ_k . It is normalized by the cluster peculiar velocity, β , to obtain the kSZ signal. In our numerical calculations below, we use the same radial profile for the ICM pressure and optical depth, which is equivalent to approximating the gas as isothermal; this is not strictly the case, but the approximation has little effect on our conclusions.

We suppose that the region is observed in several millimeter/sub-millimeter bands, producing a set of maps at N different frequencies ν_i ($i = 1, \dots, N$) that we arrange in a column vector $\mathbf{m}(x)$, a function of angular position x on the sky and whose i th component is the map $m_i(x)$ at frequency ν_i . We assume the maps are in units of thermodynamic temperature, so the CMB and kSZ signals remain constant across frequencies, and we denote the beam at frequency ν_i by b_i .

The maps contain the cluster tSZ and kSZ signals, lensed CMB anisotropies and noise,

$$\mathbf{m}(x) = y_0 \mathbf{t}_t(x) + \beta \mathbf{t}_k(x) + \mathbf{s}(x) + \mathbf{n}(x), \quad (3)$$

where $\mathbf{t}_t(x)$ is the vector whose i th component is $j_\nu(\nu_i)[b_i * T_t](x)$, the beam-convolved tSZ template modulated by the tSZ frequency spectrum, j_ν , in temperature units.

The components of the kSZ vector, $\mathbf{t}_k(\mathbf{x})$, are $[b_i * T_k](\mathbf{x})$, and those of the CMB vector, $\mathbf{s}(\mathbf{x})$, are $s_i(\mathbf{x}) = [b_i * S](\mathbf{x})$, where we denote the CMB signal on the sky as $S(\mathbf{x})$. As for Σ_m , the integration along the line-of-sight for T_l and T_k is performed out to $r = 5R_{500}$.

The unlensed (and unobservable) CMB field, $\tilde{S}(\mathbf{x})$, is transformed into the observed sky signal as (e.g., [Bartelmann & Schneider 2001](#))

$$S(\mathbf{x}) = \tilde{S}(\mathbf{x}) + \delta S(\mathbf{x}), \quad (4)$$

$$\delta S(\mathbf{x}) = \nabla S(\mathbf{x}) \cdot \nabla \phi(\mathbf{x}), \quad (5)$$

to first order in the lensing potential, which is related to the convergence,

$$\kappa(\mathbf{x}) = \Sigma(\mathbf{x})/\Sigma_{\text{cr}}, \quad (6)$$

by

$$\phi(\mathbf{x}) = \int d^2 x' \kappa(\mathbf{x}') \ln(|\mathbf{x} - \mathbf{x}'|) \quad (7)$$

$$= \phi_0 \int d^2 x' \frac{\Sigma_m(x'/\theta_m)}{\Sigma_{\text{cr}}} \ln|\mathbf{x} - \mathbf{x}'| \quad (8)$$

$$\equiv \phi_0 \Phi(x/\theta_m). \quad (9)$$

All integrals are restricted to $x' < 5\theta_{500}$. The critical surface mass density, $\Sigma_{\text{cr}} = \frac{c^2}{4\pi G} \frac{D_{\text{CMB}}}{D_{\text{clus}} D_{\text{clus-CMB}}}$, is defined in terms of the angular diameter distances D_{CMB} , D_{clus} and $D_{\text{clus-CMB}}$ between the observer and the CMB, the observer and the cluster, and the cluster and the CMB, respectively. The third equality in Eq. (9) defines our (dimensionless) model template for the lensing potential, Φ_{θ_m} , parameterized by the angular scale θ_m . It is normalized by ϕ_0 (see Eq. (2)), defined such that $\Phi_{\theta_m}(0) = 1$.

The noise term in Eq. (3), $\mathbf{n}(\mathbf{x})$, includes instrumental noise and astrophysical signals that are not related to the cluster. Examples of the latter are Galactic foreground emission, extragalactic point sources and lensing by matter randomly projected along the line-of-sight (large-scale structure, or LSS, noise). The effects of LSS correlated with cluster position can only be evaluated with numerical simulations and remain beyond the scope of the present work. Similarly, the background extragalactic point source population is modified near the cluster by lensing, creating a second order cluster-related signal that we do not consider in this work.

2.2. Reconstruction of the Lensing Potential Map

Given a map of CMB temperature anisotropy, $\hat{S}$ – obtained from a prior component separation step, as describe below – centered on a cluster, we apply the flat-sky quadratic estimator from [Hu & Okamoto \(2002\)](#) to find the Fourier modes of the projected gravitational potential:

$$\hat{\phi}(\mathbf{K}) = A(\mathbf{K}) \sum_{\mathbf{k}} \hat{S}^*(\mathbf{k}) \hat{S}(\mathbf{k}') F(\mathbf{k}, \mathbf{k}'), \quad (10)$$

with $\mathbf{K} = \mathbf{k} - \mathbf{k}' [\text{mod } n]$, where n is the number of pixels along the x (or y) axis, and

$$A(\mathbf{K}) = \left[\sum_{\mathbf{k}} f(\mathbf{k}, \mathbf{k}') F(\mathbf{k}, \mathbf{k}') \right]^{-1}. \quad (11)$$

The weights $F(\mathbf{k}, \mathbf{k}')$ are defined so that $\hat{\phi}$ is the minimum variance estimator:

$$F(\mathbf{k}, \mathbf{k}') = \frac{f^*(\mathbf{k}, \mathbf{k}')}{2P_{\hat{S}}(k)P_{\hat{S}}(k')}, \quad (12)$$

in which $P_{\hat{S}}(k) = |b(k)|^2 C_k + P_{\text{noise}}(k)$ is the observed power spectrum with a contribution from the effective noise, P_{noise} , of the cleaned CMB map; $b(k)$ represents the effective instrumental beam¹; C_k is the power spectrum of the (true) lensed CMB signal, $S(\mathbf{x})$; and $f(\mathbf{k}, \mathbf{k}')$ is given by

$$f(\mathbf{k}, \mathbf{k}') = b^*(k)b(k') [\tilde{C}_k \mathbf{k} \cdot \mathbf{K} - \tilde{C}_{k'} \mathbf{k}' \cdot \mathbf{K}], \quad (13)$$

where $\tilde{C}_k$ is the power spectrum of the unlensed CMB sky, $\tilde{S}(\mathbf{x})$.

Equation (10) gives us an unbiased estimate of $\phi(\mathbf{K})$ for a cluster of given properties in the sense that if averaged over all realizations of the (unlensed) CMB and instrumental noise, $\langle \hat{\phi}(\mathbf{K}) \rangle = \phi(\mathbf{K})$. The variance of the estimate about the mean is given by $A(\mathbf{K})$, and the reconstruction is optimal in that it minimizes this variance for each mode $\mathbf{K}$.

2.3. Matched filter

Adopting the model potential template of Eq. (9), our matched filter operates on the lensing potential map to extract the normalization, ϕ_0 , for a given scale θ_m . Each mode $\mathbf{K}$ of the estimated potential, $\hat{\phi}$, is an independently measured variable with standard deviation $A^{1/2}(\mathbf{K})$. We therefore construct the matched filter for the potential amplitude as

$$\hat{\phi}_0 = \left[\sum_{\mathbf{K}} \frac{|\Phi(\mathbf{K})|^2}{A(\mathbf{K})} \right]^{-1} \sum_{\mathbf{K}} \frac{\Phi^*(\mathbf{K})}{A(\mathbf{K})} \hat{\phi}(\mathbf{K}), \quad (14)$$

where $\Phi(\mathbf{K})$ is the Fourier transform of the model template (Eq. (9)). This yields an unbiased estimate of ϕ_0 with minimal variance given by

$$\text{Var}(\hat{\phi}_0) = \left[\sum_{\mathbf{K}} \frac{|\Phi(\mathbf{K})|^2}{A(\mathbf{K})} \right]^{-1}. \quad (15)$$

Once normalized by our measurement, $\hat{\phi}_0$, the cluster mass model is completely specified. We could quote our filter measurements directly as $\hat{\phi}_0$, but choose instead to express them in terms of the integrated convergence calculated using the model:

$$K_{5\theta_{500}} \equiv 2\pi \int_0^{5\theta_{500}} dx x \kappa(x) = \frac{1}{D_{\text{clus}}^2(z) \Sigma_{\text{cr}}(z)} M_{5R_{500}}. \quad (16)$$

Note that $\theta_{500} = R_{500}/D_{\text{clus}}(z) = c_{500}\theta_m$. The first equality defines our preferred observable, and the second relates it to cluster mass calculated within the radius $5R_{500}$; this is easily translated into M_{500} given the concentration parameter, c_{500} . Our estimator for this observable is

$$\hat{K}_{5\theta_{500}} \equiv 2\pi \hat{\phi}_0 \int_0^{5\theta_{500}} dx x \frac{\Sigma_m(x/\theta_m)}{\Sigma_{\text{cr}}(z)} \quad (17)$$

$$= \frac{1}{D_{\text{clus}}^2(z) \Sigma_{\text{cr}}(z)} \hat{M}_{5R_{500}}. \quad (18)$$

Hereafter, we consider $\hat{K}_{5\theta_{500}}$ as the output of the filter; it is dimensionless and expressed in arcmin^2 . The second line defines our cluster mass estimator, whose units are set by the prefactor.

These estimators are unbiased over the CMB and noise ensembles: $\langle \hat{K}_{5\theta_{500}} \rangle = K_{5\theta_{500}}$ and $\langle \hat{M}_{5R_{500}} \rangle = M_{5R_{500}}$. They are also optimal in that they minimize their respective variances over the same ensemble. Explicitly, we have

$$\text{Var}(\hat{M}_{5R_{500}}) = \left[D_{\text{clus}}^2(z) \Sigma_{\text{cr}}(z) \right]^2 \sigma_{K_{5\theta_{500}}}^2, \quad (19)$$

¹ Taken, for simplicity, to be axially symmetric.

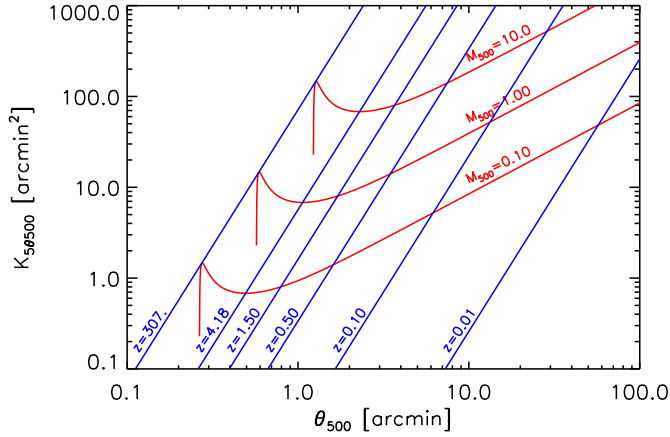


Fig. 1. Contours of cluster mass in $10^{15} M_{\odot}$ units (red curves) and redshift (blue curves) projected onto the observational plane defined by the filter output K_{500} and angular scale $\theta_{500} = c_{500}\theta_m$. A cluster of fixed mass M_{500} follows a red contour according to Eq. (16) as it moves out in redshift. Each mass contour follows the same pattern, simply displaced in amplitude. This figure shows how each point in the observational plane maps to a point in the cluster plane of (M_{500}, z) .

for the variance of the mass estimator, where $\sigma_{K_{500}}^2$ is the variance of our filter output, calculated using Eqs. (15) and (17). This is the uncertainty on the mass measurement of a single cluster, with contributions from instrumental noise and the CMB fluctuations themselves. We will see below that the kSZ adds an additional source of noise, as well as a bias term; fortunately, they are small in practice.

Figure 1 illustrates the relation between the observation plane (K_{500}, θ_{500}) and physical cluster quantities. Each point in this plane maps directly to a point in the cluster (M_{500}, z) plane, as shown by the contours of iso-mass (in red) and iso-redshift (blue). At fixed mass, Eq. (16) specifies the evolution of K_{500} as a function of redshift, which when coupled with $\theta_{500} \propto M_{500}^{1/3}/D_{\text{clus}}(z)$ describes an iso-mass curve. As the cluster moves out in redshift, its angular size decreases; at the same time, the prefactor in Eq. (16) determines the decrease and final upturn in the integrated convergence. The rapid decrease from low redshift outwards may seem surprising, but is due to the fact that we integrate over angular extent. Integrated over physical extent, the convergence increases at first with redshift as the lensing kernel becomes more efficient, but the angular extent decreases and drives down the value of K_{500} . Each mass follows the same general curve, simply displaced in absolute scale.

A given experiment will trace a sensitivity curve in the observation plane. In Fig. 2 we show predictions for the filter sensitivity, expressed by the standard deviation of the filter variance, as a function of filter angular scale for a number of different experiments. The top three curves (traced in red) all refer to space-based experiments, while the three lower curves (in blue) represent ground-based experiments similar to the three generations of SPT (Story et al. 2013; Austermann et al. 2012; Benson & Benson 2013), ACT (Das et al. 2014) and ACT-Pol (Niemack et al. 2010). The space missions are *Planck* (Planck Collaboration I 2014) and two missions proposed to the European Space Agency, CoRE (The CoRE Collaboration et al. 2011) and PRISM (PRISM Collaboration et al. 2014). The former has similar resolution to *Planck*, but more detectors and lower noise, while the latter envisions a larger telescope with both lower noise and higher angular resolution.

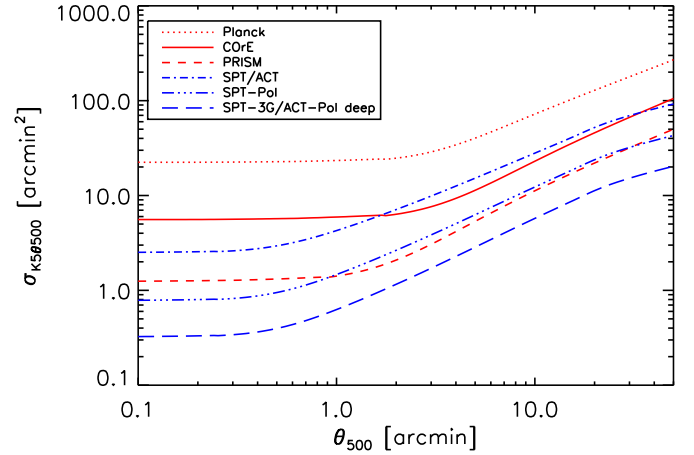


Fig. 2. Standard deviation, or filter noise, of the matched filter output (Eq. (16)) versus filter scale $\theta_{500} = c_{500}\theta_m$ for the different experimental setups, as labeled. The red curves give results for the *Planck* satellite and two future missions proposed to ESA, CoRE and PRISM. The first has similar angular resolution to *Planck* (~ 5 arcmin FWHM), but lower noise, while PRISM has both lower noise and higher resolution at ~ 2.6 arcmin. The blue curves show the noise levels for SPT/ACT, SPT-Pol and SPT-3G/ACT-Pol. All curves assume the filter is perfectly matched to the clusters. The experimental characteristics are summarized in Table 1.

We summarize the adopted characteristics of each experiment in Table 1 in terms of angular resolution and white noise in the reconstructed CMB map². We emphasize, however, that the experiments have very different frequency coverage. Although not accounted for in this present discussion, we show later that extensive frequency coverage is crucial for proper CMB reconstruction and, especially, removal of the tSZ signal.

All sensitivity curves start on small angular scales with a flat response and then break to a rise toward larger filter scales. The break occurs on smaller scales for the higher resolution ground-based experiments. CoRE and *Planck* have the same resolution and break on the same scale, but with its lower noise level, CoRE's plateau lies below that of *Planck*. PRISM has a noise level comparable to SPT-3G, but slightly lower angular resolution. We see that it breaks at an intermediate filter scale and on a higher plateau than SPT-3G. This demonstrates the interplay of angular resolution and noise: At a given noise level in $\mu\text{K}\cdot\text{arcmin}$, an experiment with higher angular resolution accesses more modes $\mathbf{k}$ to reconstruct a given potential mode $\mathbf{K}$ (Eq. (10)), thereby reducing the filter noise.

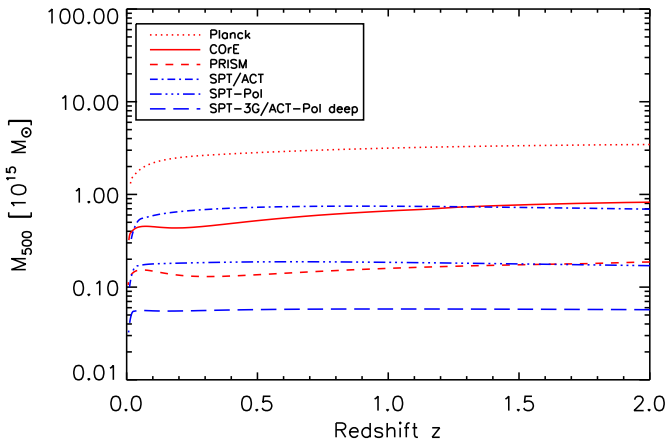
In Fig. 3 we give the standard deviation of the mass measurement, $M_{500}(z) = \sigma(z)$ (square root of Eq. (19)), as a function of redshift for each experimental setup. The results provide a useful metric for each experiment's ability to measure cluster mass: The uncertainty on the mean mass for a sample of N clusters at redshift z will be $M_{500}(z)/N^{1/2}$. Note, however, that a cluster of $nM_{500}(z)$ will have a significance smaller than $n\sigma$. This is because as the mass increases from the limiting value, the cluster's angular size also increases, driving the filter noise higher (unless we are on the small-scale plateau of the curves in Fig. 2).

We see that CoRE and SPT/ACT can only be expected to measure mass for the most massive systems in the universe, while *Planck* cannot be expected to measure any individual cluster mass. The sensitivity of these experiments, however, is

² In each case, the reconstructed CMB map was assigned the characteristics of the experimental band closest to 143 GHz.

Table 1. Characteristics of representative surveys: *Planck* (Planck Collaboration I 2014), CORE (The CORE Collaboration et al. 2011), PRISM (PRISM Collaboration et al. 2014), ACT (Das et al. 2014), ACT-Pol (Niemack et al. 2010), SPT (Story et al. 2013), SPT-Pol (Austermann et al. 2012), SPT-3G (Benson & Benson 2013).

Name	Location	Map resolution (<i>FWHM</i>) [arcmin]	Map noise [μ K-arcmin]
<i>Planck</i>	Space	5.0	45.0
CORE	Space	5.0	2.6
PRISM	Space	2.5	2.6
Generation: SPT/ACT	Ground	1.0	18.0
Generation: SPT-Pol	Ground	1.0	5.0
Generation: SPT-3G/ACT-Pol deep	Ground	1.0	2.0

**Fig. 3.** Standard deviation of mass measurements, $M_{500}(z)$, with the matched filter (square root of Eq. (19)) as a function of cluster redshift for the same experimental setups plotted in Fig. 2. This comparison does not take into account the ability of each experiment to eliminate contaminating signals, such as the tSZ, which depends on spectral coverage. In this context, the space-based missions benefit from wider spectral coverage (see text).

sufficient to obtain mean mass as a function of other cluster observables by binning measurements; in other words, to establish mean observable-mass scaling relations. Herein lies the value of our matched filter approach, by providing a means of combining many low signal-to-noise measurements to statistically determine cluster mass. The sensitivity of SPT-Pol, SPT-3G/ACT-Pol and the PRISM mission, on the other hand, is sufficient to enable individual cluster mass measurements as well as statistical studies.

3. Astrophysical contaminants

The above discussion supposes that we have a clean map of CMB anisotropies from which to extract the lensing signal. To produce this map, we must first separate the CMB component from other astrophysical emission. In this work, we focus on the potentially most troublesome signals, those produced by the cluster itself at CMB-dominant frequencies, namely the tSZ and kSZ effects.

Generally, we employ the Internal Linear Combination (ILC) methodology to separate the CMB from other signals. Because we will find that the standard ILC does not sufficiently remove the tSZ signal, causing an important bias in our mass estimation, we present two techniques based on the ILC that seek in addition to actively remove the tSZ signal: template fitting and subtraction of the tSZ, and an ILC constrained to cancel the tSZ. Both prove satisfactory.

The kSZ cannot be eliminated in this fashion, however, having the same spectral signature as the CMB. It could be removed by template fitting for experiments with sufficient resolution, but as will be seen, it does not produce any significant bias in our mass estimations. Keeping this in mind, we restrict ourselves to an examination of the nature of the kSZ contamination on the reconstructed lensing potential. This will guide our interpretation of the simulation results.

We test our methodology using the simulations described in Sect. 4, presenting the results in Sect. 5. We will see the importance of actively removing the tSZ, and that any effect from the kSZ remains manageable. Our simulations do not include diffuse Galactic or extragalactic emission not related to the clusters. At this point of our study, having employed an ILC at the heart of our component separation techniques, we assume that these are adequately controlled; the issue will be studied in more detail in future work.

3.1. The thermal SZ signal

We develop two approaches to removing the tSZ signal: multifrequency matched filters (MMF) to estimate the amplitude of the tSZ effect and then remove it from the survey frequency maps prior to a standard ILC, and an ILC constrained to eliminate the tSZ signal while extracting the CMB map. We compare the performance of the two approaches in Sect. 5. The goal is to produce the CMB map, $\hat{S}$, used as input in the lensing reconstruction of Eq. (10).

3.1.1. Multifrequency matched filters (MMF)

With this approach we apply the MMF of Melin et al. (2006) using the pressure profile $T_i(x/\theta_i)$ to obtain an estimate of the central Compton y parameter:

$$\hat{y}_0 = \int d^2x \Psi_{\theta_i}^t(x) \cdot m(x), \quad (20)$$

where

$$\Psi_{\theta_i}(k) = \sigma_{\theta_i}^2 \mathbf{P}^{-1}(k) \cdot t_i(k), \quad (21)$$

with

$$\sigma_{\theta_i} \equiv \left[\int d^2k t_i^t(k) \cdot \mathbf{P}^{-1} \cdot t_i(k) \right]^{-1/2}, \quad (22)$$

$\mathbf{P}(k)$ being the interband power spectrum matrix with contributions from (non-tSZ) sky signal and instrumental noise. It is the effective noise matrix for the MMF and can be estimated directly on the data, since the tSZ is small compared to other astrophysical signals in the sky patch centered on the cluster.

We assume that the tSZ is accurately described by our model and seek only to normalize the tSZ template to each cluster. This is the case for our simulations, because we are using the same profile for the simulated tSZ and the filter template. In reality, possible mismatch between the true cluster profile and the filter template would be a source of uncertainty.

Once we have normalized the template, we remove the tSZ signal from each of the N individual frequency maps,

$$\widehat{\mathbf{m}}(\mathbf{x}) = \mathbf{m}(\mathbf{x}) - \hat{y}_0 \mathbf{t}_t(\mathbf{x}), \quad (23)$$

and apply a standard ILC to reconstruct the clean CMB map:

$$\hat{S}(\mathbf{k}) = [\mathbf{b}' \mathbf{P}^{-1} \mathbf{b}]^{-1} \mathbf{b}' \mathbf{P}^{-1} \widehat{\mathbf{m}}(\mathbf{k}), \quad (24)$$

where $\mathbf{b}$ is the beam vector of dimension N . Finally, we convolve the resulting map with a fiducial beam, b_{fid} (5 arcmin in the case of *Planck*), to obtain the $\hat{S}(\mathbf{k})$ used in Eq. (10). The power spectrum associated with this map is

$$P_{\hat{S}} = |b_{\text{fid}}|^2 [\mathbf{b}' \mathbf{P}^{-1} \mathbf{b}]^{-1}. \quad (25)$$

3.1.2. Constrained Internal Linear Combination (CILC)

In the second approach, we directly construct the clean CMB map with a constrained ILC designed to nullify the tSZ effect, making use of its well-defined spectral signature, j_ν . The formalism is described in detail by [Remazeilles et al. \(2011\)](#). The reconstructed clean CMB map can be written

$$\hat{S}(\mathbf{k}) = \Delta^{-1} [(\mathbf{j}_\nu' \mathbf{P}^{-1} \mathbf{j}_\nu) \mathbf{b}' \mathbf{P}^{-1} - (\mathbf{b}' \mathbf{P}^{-1} \mathbf{j}_\nu) \mathbf{j}_\nu' \mathbf{P}^{-1}] \mathbf{m}(\mathbf{k}), \quad (26)$$

$$\Delta = (\mathbf{b}' \mathbf{P}^{-1} \mathbf{b}) (\mathbf{j}_\nu' \mathbf{P}^{-1} \mathbf{j}_\nu) - (\mathbf{b}' \mathbf{P}^{-1} \mathbf{j}_\nu)^2, \quad (27)$$

with $(\mathbf{j}_\nu)_i \equiv j_\nu(\nu_i) b_i$. The noise matrix, $\mathbf{P}$, is constructed as before.

We again convolve the resulting map with the fiducial beam, b_{fid} , to obtain the CMB map, $\hat{S}(\mathbf{k})$, used in Eq. (10). The power spectrum for this fiducial map is

$$P_{\hat{S}} = \frac{|b_{\text{fid}}|^2}{\Delta^2} [(\mathbf{j}_\nu' \mathbf{P}^{-1} \mathbf{j}_\nu)^2 (\mathbf{b}' \mathbf{P}^{-1} \mathbf{b}) - (\mathbf{j}_\nu' \mathbf{P}^{-1} \mathbf{j}_\nu) (\mathbf{b}' \mathbf{P}^{-1} \mathbf{j}_\nu)^2]. \quad (28)$$

The strength of the CILC removal, compared to the MMF approach, is its insensitivity to cluster modeling uncertainties (e.g., SZ profile). On the other hand, the variance of the noise in the reconstructed CMB map is higher, slightly increasing the error on the final mass estimate, as we show in Sect. 5. We also present the impact of tSZ modeling uncertainties on the MMF approach in the same section.

3.2. The kinetic SZ signal

After the tSZ, the kSZ is the dominant cluster signal at CMB frequencies. It has the same spectrum as the primary CMB and cannot be removed by spectral separation methods.

We can appreciate the effect of the kSZ by returning to the lensing potential reconstruction of Eq. (10). Even in the best possible case, the reconstructed CMB map, $\hat{S}$, contains kSZ in addition to the lensed CMB and noise: $\hat{S}(\mathbf{x}) = S(\mathbf{x}) + \beta T_k(\mathbf{x}/\theta_k) + n(\mathbf{x})$. The lensing map reconstruction therefore has

four contributions:

$$\hat{\phi}'(\mathbf{K}) = \hat{\phi}(\mathbf{K}) + 2\mathcal{R} \left[\beta A(\mathbf{K}) \sum_{\mathbf{k}} S(\mathbf{k}) T_k^*(\mathbf{k}'\theta_k) F(\mathbf{k}, \mathbf{k}') \right] \quad (29)$$

$$+ 2\mathcal{R} \left[\beta A(\mathbf{K}) \sum_{\mathbf{k}} n(\mathbf{k}) T_k^*(\mathbf{k}'\theta_k) F(\mathbf{k}, \mathbf{k}') \right] \quad (30)$$

$$+ \beta^2 A(\mathbf{K}) \sum_{\mathbf{k}} T_k(\mathbf{k}\theta_k) T_k^*(\mathbf{k}'\theta_k) F(\mathbf{k}, \mathbf{k}') \quad (31)$$

$$\equiv \hat{\phi}(\mathbf{K}) + \beta [\hat{\phi}_{\text{ks}}(\mathbf{K}) + \hat{\phi}_{\text{kn}}(\mathbf{K})] + \beta^2 \hat{\phi}_{\text{kk}}(\mathbf{K}) \quad (32)$$

where $\hat{\phi}$ contains just the CMB and noise terms, as considered in Sect. 2. The kSZ adds cross terms of the kSZ with both CMB, $\hat{\phi}_{\text{ks}}$, and noise, $\hat{\phi}_{\text{kn}}$, and a term quadratic in the kSZ signal, $\hat{\phi}_{\text{kk}}$.

Averaged over the CMB and noise ensembles, the two middle terms vanish, independent of the value of β , because $\langle s(\mathbf{k}) \rangle = \langle n(\mathbf{k}) \rangle = 0$. They act as an additional noise contribution to the potential reconstruction for a given cluster. This behaviour differs from that of the kSZ when directly using the temperature anisotropy induced by cluster lensing, rather than reconstruction of the lensing potential as done here. For a given β , the kSZ is guaranteed to contribute to the temperature anisotropy at a level comparable to the lensing signal ([Seljak & Zaldarriaga 2000](#); [Lewis & Challinor 2006](#)); in our case, however, the contribution could be small, depending only on the chance alignment of CMB and kSZ modes. In either case, additional averaging over a set of clusters will further reduce the effect of these terms linear in β because the objects will have random peculiar velocities.

The last term, quadratic in β , is a bias. Its presence is independent of the CMB and noise ensembles, and it cannot be beaten down by averaging over a cluster ensemble. With proper modeling of the kSZ signal, the bias could be eliminated, if needed, through subtraction, cluster by cluster, to leave a zero-mean residual as a noise contribution. This will not prove necessary in our subsequent study, where we will find that the bias is unimportant for realistic cluster velocities.

4. Simulations

We illustrate our mass estimation technique through recovery of the mass scale for a sample of clusters with simulated *Planck*-like observations. We proceed by first simulating a sample of identical clusters, and then consider a mock of the ESZ-XMM, a subsample of 62 clusters from the *Planck* Early Sunyaev-Zeldovich list (ESZ, [Planck Collaboration VIII 2011](#)) with good X-ray observations, including X-ray determined masses, M_{500}^X , spanning the range $[2-20] \times 10^{14} M_\odot$ (ESZ-XMM, [Planck Collaboration XI 2011](#)). In a companion paper, we report an estimation of the mass scale of the actual ESZ-XMM sample using the *Planck* dataset.

Our first simulation consists of 62 observations of a mock of A2163, assigning mass and tSZ profiles following our adopted templates. With an X-ray deduced mass of $M_{500}^X = 1.9 \times 10^{15} M_\odot$ ($z = 0.203$), this system is one of the most massive clusters known, falling near the 1σ curve for *Planck* shown in Fig. 3, and the most massive member of the ESZ-XMM.

We generate 62 independent realizations of primary CMB anisotropies and of white noise in tangential map projections of $10 \times 10 \text{ deg}^2$ across the six highest frequency channels (100–857 GHz) of *Planck*. The mock of A2163 is centered in each channel map, the lensing is applied to the primary CMB anisotropies and the tSZ signal added. We then smooth each channel map by its corresponding beam and add the white

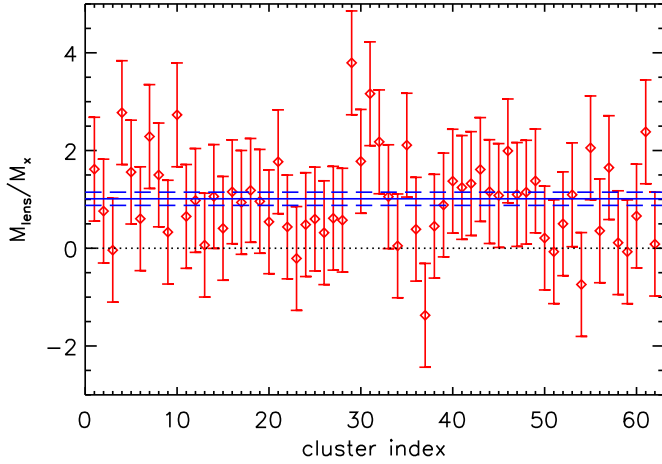


Fig. 4. Recovered mass for the 62 *Planck* simulations of A2163, expressed as the ratio of the CMB-measured mass to the input mass from the X-ray model, $\hat{M}_{500}/M_{500}^X$. The tSZ was removed in this example using the CILC. Each diamond is the result of a single analysis chain (simulation, CILC, lensing extraction and matched filtering) accompanied by its 1σ uncertainty of 1.06. The solid blue line shows the sample average (1.01) and its 1σ range ($\pm 1.06/\sqrt{62} = \pm 0.13$).

noise, taking the noise and beam characteristics as published in *Planck Collaboration I* (2014). A simulated cluster observation thus comprises six channel maps, and there are 62 such simulated observations.

For each of these mock observations, we first remove the tSZ signal and produce a clean CMB map as described above. We then reconstruct the lensing potential map and apply the matched filter to extract our mass estimate, $\hat{M}_{500}$. Note that the filter is perfectly matched to the cluster in that the filter template and actual cluster projected potential are identical. We refer to the complete processing of a single cluster observation as an analysis *chain*.

For each chain, we compare the mass measurement to the input (X-ray deduced) mass of A2163 by forming the ratio $\hat{M}_{500}/M_{500}^X$, and then take the sample mean, $\langle \hat{M}_{500}/M_{500}^X \rangle$, over the 62 observations. An unbiased recovery of the sample mass scale corresponds to $\langle \hat{M}_{500}/M_{500}^X \rangle = 1$.

We run two sets of 62 chains without kSZ to compare the results from the two different component separation methods presented in Sect. 3. We also ran additional simulation chains adding the kSZ effect with constant (systematic) peculiar velocities of 300 km s^{-1} and 900 km s^{-1} to each of the 62 clusters.

5. Results

Figure 4 shows the recovered mass ratio $\hat{M}_{500}/M_{500}^X$ for the 62 simulated cluster observations without kSZ when using the constrained ILC to remove the tSZ. The individual measurement uncertainty on this quantity is large at 1.06 (the same for all chains, because the cluster and the statistical CMB and noise properties are the same for the 62 realizations). The points clearly disperse preferentially above zero, and taking the sample average we find $\langle \hat{M}_{500}/M_{500}^X \rangle = 1.01 \pm 0.13$.

We have an unbiased recovery of the sample mass scale with 13% uncertainty. The result is identical when using the MMF component separation procedure³. The basic result of this study,

³ The individual measurement uncertainty is slightly lower in this case (1.04 against 1.06), but the sample mean and its uncertainty are the same to the given precision.

therefore, is that we detect the mass scale of the sample at greater than 7σ .

If we model the SZ emission with the cool-core profile given in Table C.2. of *Arnaud et al. (2010)* but extract it with the MMF based on the universal profile, the sample average increases from 1.01 ± 0.13 to 1.19 ± 0.13 . Adopting instead the morphologically disturbed profile, the value shifts to 1.09 ± 0.13 . The mis-modeling of the SZ profile in the MMF component separation procedure may thus introduce a bias of order $\sim 1\sigma$ on the average.

We have also tested the sensitivity of our conclusions to the extension of the mass profile by truncating the integration at R_{vir} instead of $5R_{500}$. The individual measurement uncertainty increases to 1.23, which in turn slightly increases the uncertainty on the sample average to 0.16. This results in a modest decrease in the global significance of the detection from 7σ to between 6 and 7σ .

Accurate removal of the tSZ signal is essential, something which can be gauged from the results when applying a standard ILC to extract the CMB without any constraint to nullify the tSZ. In this case, we find a sample mean of $\langle \hat{M}_{500}/M_{500}^X \rangle = 1.93 \pm 0.13$, highly biased by the residual tSZ signal.

The standard ILC is incapable of removing the tSZ signal to a sufficiently high level. We expect that this is in large part due to the fact that the tSZ is only a weak component in the map and hence not accounted for by the standard ILC weights. Our component separation techniques manage to adequately remove the tSZ by direct subtraction (MMF) or cancelation (CILC), both relying on the known spectral dependence of the signal. It is clear that multi-band CMB observations for accurate removal of the tSZ are an important consideration in designing experimental campaigns.

Turning to the kSZ effect, we find a sample mean of $\langle \hat{M}_{500}/M_{500}^X \rangle = 1.00 \pm 0.13$ for the case of 300 km s^{-1} constant peculiar velocity, and $\langle \hat{M}_{500}/M_{500}^X \rangle = 0.86 \pm 0.13$ for the case of 900 km s^{-1} . This is for the CILC, but the results are essentially the same for the MMF. Note that the uncertainties on the sample means are unchanged, because they are calculated from the sample size and the individual measurement error, the latter determined by the unchanged CMB and noise properties.

There is no evidence of bias in the sample mean at the lower peculiar velocity of 300 km s^{-1} , while a bias of 14% appears at the higher value of 900 km s^{-1} . In the standard Λ CDM model, we expect individual cluster peculiar velocities to follow a Gaussian distribution of zero mean and variance $\langle \beta^2 \rangle \approx (300 \text{ km s}^{-1})^2$. Our result for the mean of a set of clusters with this constant velocity is therefore representative of the bias expected of the quadratic term in β . The bias term is clearly present, but only important at atypically large peculiar velocities.

The linear term is also present, causing an increase in the observed dispersion of the individual measurements. By comparing the dispersion before and after addition of the kSZ effect, we deduce that it contributes 0.14 and 0.31, respectively, for the lower and higher peculiar velocities. The former value is the more realistic and should be compared to the CMB and noise contribution to the dispersion (individual uncertainties) of 1.06.

The overall conclusion is the same for the simulated ESZ-XMM, as shown in Fig. 5. We obtain an unbiased estimate of the sample mass scale: $\langle \hat{M}_{500}/M_{500}^X \rangle = 0.99 \pm 0.28$. The significance is lower than before because the ESZ-XMM contains a range of cluster masses, all of them smaller than A2163, its most massive member. Nevertheless, we find that the mass scale can be recovered at the 3σ level. In this case, truncating the profile

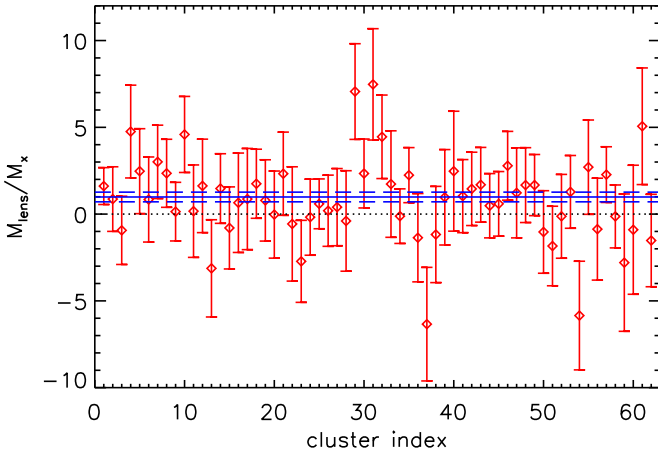


Fig. 5. Recovered mass for the *Planck* ESZ-XMM simulation, expressed as in Fig. 4 by the ratio of the CMB-measured mass to the input mass from the X-ray model, $\hat{M}_{500}/M_{500}^x$. The tSZ was removed using the CILC, and the blue solid and dashed lines show the sample mean and its 1σ range: $\langle \hat{M}_{500}/M_{500}^x \rangle = 0.99 \pm 0.28$. This simulation includes a random kSZ effect with Gaussian standard deviation of 300 km s^{-1} . The CMB and noise realizations are the same as in Fig. 4, resulting in a similar pattern in the distribution of points, but individual uncertainties are larger (note the change in scale).

at R_{vir} increases the uncertainty from 0.28 to 0.33, leaving the conclusion unchanged. This simulation also included a kSZ effect from random peculiar velocities with a Gaussian standard deviation of 300 km s^{-1} . We see no evidence for its impact on the recovered sample mean.

6. Discussion

Our results show that the proposed method can be a practical tool for estimating cluster masses, even with current CMB temperature data. It offers a new way to constrain cluster scaling relations between total mass and observables such as X-ray luminosity, SZ signal strength or richness. These relations are central to cosmological and large-scale structure analyses of cluster catalogues.

The method complements others for measuring cluster mass. Like gravitational shear, it directly probes total mass without assumptions about the state of any particular cluster component; this is strictly true for the CILC method, although we note that there does remain some modeling uncertainty when applying the MMF component separation, as discussed above. The method's particular strengths are that it can be used to much higher redshifts, using the CMB as a source plane, and that it is sensitive to the convergence field, rather than its gradient like the shear. Noting this latter difference, several authors have recently pointed out the value of combined CMB and shear analyses in the more general context (Hand et al. 2013; Das et al. 2013).

A critical capability of the proposed procedure is to provide accurate measurements and their uncertainties in the low signal-to-noise regime. This enables statistical analyses that permit practical application to existing CMB temperature datasets.

We achieve it through the lensing reconstruction that furnishes well-defined noise properties for use by the filter. The noise arises not just from the instrument, but also from the CMB itself, because the lensing signature is a correlation between CMB anisotropy modes behind the cluster, modes that we do not a priori know. Instead, we rely on the power spectrum of the primary CMB (and its Gaussianity) to tell us what they are

on average and the dispersion about that average. In many ways, this is simpler than trying to determine the CMB gradient around each individual cluster, as needed when working directly with the secondary temperature anisotropy generated by lensing.

Figure 3 summarizes the ideal statistical power of various experiments to measure cluster mass, showing the standard deviation of filter mass determinations as a function of cluster redshift. It is ideal because it only accounts for noise from the CMB and the instrument. The curves are extremely flat in redshift, a reflection of the broad lensing kernel to the source plane of the CMB.

Comparing the space experiments (the red curves), we see that *Planck* is dominated by instrumental noise, since it has the same angular resolution as CORe, which performs much better with its lower noise level, while PRISM gains further by incorporating more modes in reaching to smaller scales. Even at their higher angular resolution, the ground-based experiments are not dominated by the CMB, as can be seen from the fact that the mass filter noise continues to decrease with decreasing instrumental noise.

We examined the impact of the tSZ and kSZ signals using simulations of *Planck*-like observations of massive clusters. Our main result is that the tSZ must be accurately removed to avoid biasing the mass estimations. A standard ILC is not sufficient. We applied two methods, both of which proved satisfactory. Based on the ILC, their key additional characteristic is that they actively subtract or nullify the tSZ. The success of the CILC is encouraging, because it does not rely on any adopted profile for the tSZ. Our result emphasizes the importance of broad frequency coverage in experimental design to enable adequate component separation.

The kSZ cannot be removed through such spectral separation methods and it remains in the CMB maps used to reconstruct the lensing potential. It impacts the final result by adding a source of measurement noise (term linear in β) and a bias (term in β^2). In our simulations, we found evidence for both terms. Fortunately, their influence is small for the expected distribution of cluster peculiar velocity.

These latter conclusions only apply in the context of our simulations of *Planck*-like CMB observations of massive clusters. Understanding these details in other experimental setups, e.g., ground-based instruments, would require dedicated simulations. The same applies to study of *Planck*-like observations of large cluster samples including lower mass systems.

Additional limitations of the simulations presented here include lack of other foreground sources, such as diffuse Galactic emission and extragalactic sources. Since we are using component separation techniques based on the ILC, we assume for this preliminary study that they adequately remove these foregrounds. Our clusters are also simulated in isolation and modeled by the same spherically symmetric profiles used in our filters (tSZ and mass filters)⁴. More realistic simulations would employ variations in model profiles and numerical simulations of clusters in their cosmological context to evaluate the effects of local structure around the clusters, as well as that of uncorrelated large-scale structure along the line-of-sight that contribute to the noise term.

Despite these limitations, our simulations are sufficient to demonstrate the potential of the *Planck* 2013 dataset to detect the mass scale of the ESZ-XMM. Our simulation of this observation of 62 clusters is summarized in Fig. 5. The conclusion is

⁴ Recall the above statement that the CILC does not, however, rely on an adopted template.

that we should be able to detect the mass scale of this catalog to slightly more than 3σ . A separate paper presents our analysis on the real *Planck* dataset.

7. Conclusion

We propose a method to measure galaxy cluster masses using CMB lensing and demonstrate that it can be practically applied to existing datasets (e.g., ACT, *Planck*, SPT) in statistical analyses of cluster samples. The strength of the approach draws from its ability to provide viable mass estimates and uncertainties even in low signal-to-noise regimes, thereby enabling straightforward statistical analyses of systems well below individual detection.

Accurate removal of the tSZ is important and achievable, as we demonstrate by application of component separation methods that actively subtract or nullify it. The implication is that experimental design must allow for sufficient spectral coverage to enable effective separation methods.

Using a simulation of the *Planck* ESZ-XMM sample, we conclude that it would be possible to determine the mass scale of this set of 62 clusters to 3σ significance (CMB and instrumental noise only). In a companion paper, we present a first application of our method to the *Planck* data on the actual ESZ-XMM.

The method presented here uses only temperature data in the lensing reconstruction. Future work will extend it to CMB polarization data. Our preliminary study here opens the way to numerous research avenues targeting additional issues related to foregrounds and large-scale structure, and calls for detailed studies dedicated to specific experimental campaigns.

Lensing of the CMB opens a new and independent avenue for studying cluster masses, an important complement to other techniques, such as weak gravitational lensing of background galaxies. In fact, CMB lensing offers the possibility of calibrating large cluster samples now while we await large area galaxy lensing surveys, such as the Dark Energy Survey, the Large Synoptic Survey Telescope, and the *Euclid* and WFIRST space missions. And it will remain the more efficient way to measure cluster masses at high redshifts, where the source galaxy population rapidly declines in imaging surveys.

Acknowledgements. The authors would like to thank the anonymous referee for useful comments which helped to clarify some important aspects of this work. A portion of the research described in this paper was carried out at the Jet Propulsion Laboratory, California Institute of Technology, under a contract with the National Aeronautics and Space Administration.

References

Allen, S. W., Evrard, A. E., & Mantz, A. B. 2011, *ARA&A*, 49, 409
 Arnaud, M., Pointecouteau, E., & Pratt, G. W. 2005, *A&A*, 441, 893
 Arnaud, M., Pointecouteau, E., & Pratt, G. W. 2007, *A&A*, 474, L37
 Arnaud, M., Pratt, G. W., Piffaretti, R., et al. 2010, *A&A*, 517, A92
 Austermann, J. E., Aird, K. A., Beall, J. A., et al. 2012, in *SPIE Conf. Ser.*, 8452

Bartelmann, M., & Schneider, P. 2001, *Phys. Rep.*, 340, 291
 Benson, B., & Benson. 2013, in *IAU Symp.* 288, eds. M. G. Burton, X. Cui, & N. F. H. Tothill, 76
 Bocquet, S., Saro, A., Mohr, J. J., et al. 2015, *ApJ*, 799, 214
 Böhringer, H., Chon, G., & Collins, C. A. 2014, *A&A*, 570, A31
 Bullock, J. S., Kolatt, T. S., Sigad, Y., et al. 2001, *MNRAS*, 321, 559
 Das, S., Sherwin, B. D., Aguirre, P., et al. 2011, *Phys. Rev. Lett.*, 107, 021301
 Das, S., Errard, J., & Spergel, D. 2013, *ArXiv e-prints* [[arXiv:1311.2338](#)]
 Das, S., Louis, T., Nolte, M. R., et al. 2014, *J. Cosmol. Astropart. Phys.*, 4, 14
 Gladders, M. D., Yee, H. K. C., Majumdar, S., et al. 2007, *ApJ*, 655, 128
 Hand, N., Leauthaud, A., Das, S., et al. 2013, *ArXiv e-prints* [[arXiv:1311.6200](#)]
 Hasselfield, M., Hilton, M., Marriage, T. A., et al. 2013, *Cosmol. Astropart. Phys.*, 7, 8
 Hoekstra, H., Herbonnet, R., Muzzin, A., et al. 2015, *MNRAS*, 449, 685
 Holder, G., & Kosowsky, A. 2004, *ApJ*, 616, 8
 Hu, W., & Okamoto, T. 2002, *ApJ*, 574, 566
 Lewis, A., & Challinor, A. 2006, *Phys. Rep.*, 429, 1
 Mantz, A., Allen, S. W., Rapetti, D., & Ebeling, H. 2010, *MNRAS*, 406, 1759
 Mantz, A. B., von der Linden, A., Allen, S. W., et al. 2015, *MNRAS*, 446, 2205
 Marrone, D. P., Smith, G. P., Okabe, N., et al. 2012, *ApJ*, 754, 119
 Maturi, M., Bartelmann, M., Meneghetti, M., & Moscardini, L. 2005, *A&A*, 436, 37
 Melin, J.-B., Bartlett, J. G., & Delabrouille, J. 2006, *A&A*, 459, 341
 Muñoz-Cuartas, J. C., Macciò, A. V., Gottlöber, S., & Dutton, A. A. 2011, *MNRAS*, 411, 584
 Nagai, D., Kravtsov, A. V., & Vikhlinin, A. 2007, *ApJ*, 668, 1
 Navarro, J. F., Frenk, C. S., & White, S. D. M. 1996, *ApJ*, 462, 563
 Neto, A. F., Gao, L., Bett, P., et al. 2007, *MNRAS*, 381, 1450
 Niemack, M. D., Ade, P. A. R., Aguirre, J., et al. 2010, in *SPIE Conf. Ser.*, 7741
 Planck Collaboration VIII. 2011, *A&A*, 536, A8
 Planck Collaboration XI. 2011, *A&A*, 536, A11
 Planck Collaboration I. 2014, *A&A*, 571, A1
 Planck Collaboration XVI. 2014, *A&A*, 571, A16
 Planck Collaboration XVII. 2014, *A&A*, 571, A17
 Planck Collaboration XX. 2014, *A&A*, 571, A20
 Planck Collaboration 2015, *A&A*, submitted [[arXiv:1502.01597](#)]
 Planck Collaboration Int. III. 2013, *A&A*, 550, A129
 Pratt, G. W., Croston, J. H., Arnaud, M., & Böhringer, H. 2009, *A&A*, 498, 361
 PRISM Collaboration 2014, *JCAP*, 2, 6
 Reichardt, C. L., Stalder, B., Bleem, L. E., et al. 2013, *ApJ*, 763, 127
 Remazeilles, M., Delabrouille, J., & Cardoso, J.-F. 2011, *MNRAS*, 410, 2481
 Rozo, E., Rykoff, E. S., Evrard, A., et al. 2009, *ApJ*, 699, 768
 Rozo, E., Wechsler, R. H., Rykoff, E. S., et al. 2010, *ApJ*, 708, 645
 Rozo, E., Rykoff, E. S., Bartlett, J. G., & Evrard, A. E. 2013, *ArXiv e-prints* [[arXiv:1302.5086](#)]
 Rozo, E., Bartlett, J. G., Evrard, A. E., & Rykoff, E. S. 2014a, *MNRAS*, 438, 78
 Rozo, E., Evrard, A. E., Rykoff, E. S., & Bartlett, J. G. 2014b, *MNRAS*, 438, 62
 Seljak, U., & Zaldarriaga, M. 2000, *ApJ*, 538, 57
 Stanek, R., Evrard, A. E., Böhringer, H., Schuecker, P., & Nord, B. 2006, *ApJ*, 648, 956
 Story, K. T., Reichardt, C. L., Hou, Z., et al. 2013, *ApJ*, 779, 86
 Sunyaev, R. A., & Zeldovich, Y. B. 1972, *Comm. Astrophys. Space Phys.*, 4, 173
 Sunyaev, R. A., & Zeldovich, I. B. 1980, *MNRAS*, 190, 413
 The CoRE Collaboration, Armitage-Caplan, C., Avillez, M., et al. 2011, *ArXiv e-prints* [[arXiv:1102.2181](#)]
 Vale, C., Amblard, A., & White, M. 2004, *New Astron.*, 10, 1
 van Engelen, A., Keisler, R., Zahn, O., et al. 2012, *ApJ*, 756, 142
 Vikhlinin, A., Kravtsov, A., Forman, W., et al. 2006, *ApJ*, 640, 691
 Vikhlinin, A., Kravtsov, A. V., Burenin, R. A., et al. 2009, *ApJ*, 692, 1060
 von der Linden, A., Mantz, A., Allen, S. W., et al. 2014, *MNRAS*, 443, 1973
 Yee, H. K. C., & Ellingson, E. 2003, *ApJ*, 585, 215
 Zaldarriaga, M., & Seljak, U. 1999, *Phys. Rev. D*, 59, 123507