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HAL Id: cea-01135401
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Submitted on 25 Mar 2015

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Determination of the far-infrared dust opacity in a prestellar core

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Received 23 February 2012 / Accepted 11 June 2013

ABSTRACT

Context. Mass estimates of interstellar clouds from far-infrared and submillimetre mappings depend on the assumed dust absorption cross-section for radiation at those wavelengths.

Aims. The aim is to determine the far-IR dust absorption cross-section in a starless, dense core located in Corona Australis. The value is needed for determining of the core mass and other physical properties. It can also have a bearing on the evolutionary stage of the core.

Methods. We correlated near-infrared stellar H − Ks colour excesses of background stars from NTT/SOFI with the far-IR optical depth map, from Herschel observations at four wavelengths. In this way we correlated near-infrared photometry in the Ks bands in conjunction with Herschel maps at four wavelengths. In this way we correlated near-infrared photometry in the Ks bands in conjunction with Herschel maps at four wavelengths.

Results. A linear correlation is seen between the colour H − Ks and τFR up to high extinctions (AV ∼ 25). The correlation translates to the average extinction ratio A250/μm/Av = 0.0014 ± 0.0002, assuming a standard near-infrared extinction law and a dust emissivity index β = 2. Using an empirical NHI/Av ratio we obtain an average absorption cross-section per H nucleus of σ250=μm = (1.8±0.3)×10^{-25} cm² H-atom^{-1}, corresponding to a cross-section per unit mass of gas k250=μm = 0.08±0.01 cm² g⁻¹. The cloud model, however, suggests that owing to the bias caused by temperature changes along the line-of-sight, these values underestimate the true cross-sections by up to 40% near the centre of the core. Assuming that the model describes the effect of the temperature variation on τFR correctly, we find that the relationship between H − Ks and τFR agrees with the recently determined relationship between σ250 and NHI in Orion A.

Conclusions. The derived far-IR cross-section agrees with previous determinations in molecular clouds with moderate column densities, and is not particularly large compared with some other cold cores. We suggest that this is connected to the core not being very dense (the central density is likely to be ~10^5 cm⁻³), and judging from previous molecular line data, it appears to be at an early stage of chemical evolution.

Key words. ISM: clouds – dust, extinction

1. Introduction

Determining the emission properties of interstellar dust is useful not only for providing reliable molecular cloud mass estimates from thermal dust emission maps, but also for testing ideas about dust evolution. In several previous studies the far-infrared to visual or far-IR to near-IR extinction ratio (Av/Av or A2/Av), sometimes called the emissivity, has been studied by combining near-infrared photometry and dust continuum maps (Bianchi et al. 2003; Kramer et al. 2003; Shirley et al. 2005; Lehtinen et al. 2007; Shirley et al. 2011). One of the results from these studies is that the dust emissivity changes from region to region, and it shows a tendency to increase with decreasing dust temperature (Arab et al. 2012). This trend agrees well with theoretical predictions, is thought to be caused primarily by dust coagulation in the cold, dense cores of molecular clouds, and is further accentuated by ice formation on the surfaces of dust grains

* Based partly on observations collected at the European Southern Observatory, Chile (ESO programmes 075.C-0748 and 077.C-0562).
2. On JHK photometry at high extinctions

Analysis of JHK near-infrared photometry involves some uncertainties and complications. The JHK system is not uniquely defined because different detectors and filter sets have been used. Additionally, the transmissions in the JHK bands depend strongly on the atmospheric transmission (atmospheric H$_2$O absorption lines) and thus vary depending on the observing site and is even subject to night-to-night variations. As a consequence one should be careful when comparing photometric data obtained at different epochs and observing sites. This is also true when adapting a JHK reddening slope because these slopes depend on the particular filter-detector combination. The modified $K_s$ is narrower than the $K$ filter and excludes some of the strong atmospheric lines and is thus less influenced by the atmospheric opacity.

The sensitivity of near-infrared observations has increased dramatically during the last decade, and it is now possible to observe routinely highly reddened, faint stars. The reddening affects the observed object spectral energy distribution (SED) and for highly reddened sources the effective wavelength of the filter shifts towards longer wavelengths. This introduces a new effect as the JHK reddening slope changes depending on the reddening. This phenomenon was already noted in the UBV system (see e.g. Golay (1974)), and has been later studied e.g. using the Two Micron All Sky Survey (2MASS) (Straižys & Lazauskaitė 2008), and the UKIRT Infrared Deep Sky Survey (UKIDSS) data (Stead & Hoare 2009). As noted in Stead & Hoare (2009) one should refer to a reddening track instead of a reddening slope. The importance of the photometric system is highlighted by the fact that in the 2MASS system the reddening track of highly reddened giants curves down in the JHK colour–colour diagram whereas in the MKO system (UKIDSS) the track curves up. In addition to reddening this effect depends on the spectral type of the source. The effect is noticeable for $H-K$ indices larger than ~1.5. In the present work the highest $H-K_s$ indices are nearly 3.

3. Data acquisition

3.1. NTT/SOFI observations

The dense core region of CrA C was imaged in $J$, $H$, and $K_s$ using the Son Of Isaac (SOFI) near-infrared instrument on the New Technology Telescope (NTT) at the La Silla observatory in May 2005 and July 2006. The SOFI field of view is 4.9'×9.2'

1.http://irsa.ipac.caltech.edu/applications/planck/

The total on source time was 40 min, 104 min and 540 min in the $J$, $H$, and $K_s$ filters, respectively. The average seeing was ~0'.7.

The IRAF$^3$ external package XDIMSUM was used to reduce the imaging data. The images were searched for cosmic rays and sky-subtracted. The two nearest background images in time to each image were used in the sky subtraction. An object mask was constructed for each image. Applying these masks in the sky subtraction produced hole masks for each sky-subtracted image. Special dome flats and illumination correction frames provided by the NTT team were used to flat field and to illumination correct the sky-subtracted SOFI images. Rejection masks combined from a bad pixel mask and individual cosmic ray and hole masks were used when averaging the registered images. The coordinates of the registered images were derived from the 2MASS catalogue (Skrutskie et al. 2006).

The SExtractor software version 2.5.0 (Bertin & Arnouts 1996) package was used to obtain stellar photometry of the reduced SOFI images. Galaxies were excluded using the SExtractor CLASS keyword and by visual inspection of the images.

The zero points of the coadded data were fixed using standard star measurements and checked using stars in common with the less deep HAWK-I imaging of CrA reported in Sicilia-Aguilar et al. (2011) (see below). The limiting magnitudes (for a formal error of 0.1 mag) are approximately 21$^m$7, 20$^m$5, and 20$^m$8 for $J$, $H$, and $K_s$ respectively. After elimination of the nonstellar sources 25, 266, and 369 stars remained in $J$, $H$, and $K_s$ respectively. In the subsequent analysis we use the 266 stars detected both in the $H$ and $K_s$ filters. The locations of these stars are indicated with blue plus signs in Fig. 4. As one can see in this Figure, no stars were detected in the very centre of CrA C. All detections lie at the edges of the core.

The extracted $JHK_s$ magnitudes were converted into Persson magnitudes, and from these to the 2MASS photometric system as described in Ascenso et al. (2007). However, these conversion formulae could not be directly applied to all of the SOFI data because for most stars $J$ band magnitudes were not available owing to the high extinction. Using the Eqs. (1)–(6) of Ascenso et al. (2007) the conversion formula for the $(H-K_s)$ index from the instrumental magnitudes to the 2MASS system can be written as

$$(H-K_s)_{2mass} = 1.019(H-K_s) - 0.046(J-K_s) + 0.005,$$  

where the magnitudes on the right are SOFI magnitudes. We have attempted to decrease the systematic error arising from the missing $J$ magnitude by estimating the $(J-K_s)$ colour using the observed $H-K_s$ index, the intrinsic $J$ and $H-K_s$ colours of giant stars from Bessell & Brett (1988), and the extinction ratios $A(I)/A(K_s)$ derived by Indebetouw et al. (2005). These give

$$(J-K_s)_{est} = (J-K_s) - \frac{E(J-K_s)}{E(H-K_s)}(H-K_s) + \frac{E(J-K_s)}{E(H-K_s)}(H-K_s)$$  

$$= (0.39 \pm 0.11) + (2.72 \pm 0.48) \times (H-K_s).$$
The ratio of colour excesses \( E(J - K_s)/E(H - K_s) = 2.72 \pm 0.48 \) is obtained using the average \( A(J)/A(K_s) \) and \( A(H)/A(K_s) \) values listed in Table 1 of Indebetouw et al. (2005). The corresponding values of \( (J - K_0) - \frac{E(J-K)}{E(H-K)}(H - K_0) \) range from 0.27 to 0.50 for giants of the types G0 to M7 using the intrinsic colours given in Bessell & Brett (1988). The possible variation of the intrinsic colours combined with the uncertainty of the colour excess ratio causes an additional uncertainty of \( \sim 0.03 \) to the derived \( H - K_s \) index. This small uncertainty is added in quadrature to the uncertainties of \( H - K_s \) owing to formal photometric errors (\( \sim 0.09 \)) and the uncertainty in the magnitude zero-point (\( \sim 0.07 \) for \( H - K_s \)).

The conversion to the 2MASS system using Eqs. (1) and (2) reduces the instrumental \( H - K_s \) indices by about 10 percent. The observed \( H - K_s \) indices for the highly reddened stars in our sample range from \( 1^m5 \) to \( 2^m9 \), and for these stars the correction lies between \( -0^m14 \) and \( -0^m25 \). Ignoring the colour-dependent term in the conversion to the 2MASS system would bias systematically the \( H - K_s \) indices used in the following analysis.

### 3.2. VLT/HAWK-I observations

The core CrA C was partially covered by two \( 7^m5 \times 7^m5 \) fields taken with the HAWK-I IR camera on the VLT as part of the survey of Corona Australis reported in Sicilia-Aguilar et al. (2011) (ESO program 083.C-0079). A more detailed description of the observations and reduction procedures can be found in that paper. The photometry zero point of the HAWK-I data was fixed to 2MASS objects present in the fields. The typical calibration errors range from 1% to 5%, depending on the weather conditions. Therefore, signal to noise and background remain the main sources of uncertainty, especially for the fainter objects. The data are complete in the following dynamical ranges: \( J = 11 - 18.5 \) mag, \( H = 11 - 18.5 \) mag, \( K = 10.5 - 18 \) mag. The locations of HAWK-I stars with reasonably small errors (<0\(^\prime\)1 in \( H \) and \( K \)) lying within a radial distance of 5\(^\prime\)5 from the core centre are marked with red crosses in Fig. 4. The total number of these stars is 964.

Of the 266 stars in our SOFI sample 109 were detected with HAWK-I. In Fig. 1 we correlate the \( H - K_s \) indices of stars common to both surveys. Only a zero-point correction to the 2MASS system has been done to the magnitudes presented in this figure. This is because no colour-dependent correction has been applied to the HAWK-I data (Aurora Sicilia-Aguilar, priv. comm.) as the characteristics of the HAWK-I filters needed for this correction are not available. One can see the uncorrected SOFI colours agree with the HAWK-I data. After the colour correction for SOFI, the slope of the correlation deviates clearly from unity because this correction reduces the \( H - K_s \) indices by about 10%. Because our main purpose is to derive the dust opacity towards the core where the colour correction has a noticeable effect, we have decided not to use the HAWK-I data in the subsequent analysis.

### 3.3. 2MASS data

Of the stars detected in the 2MASS survey 58 (again, including only those with \( \Delta K < 0^m1 \)) lie within a distance of 5\(^\prime\)5 from the core centre. The positions of these stars (excluding those coincident with HAWK-I or SOFI stars) are marked in Fig. 4. Only two of the 2MASS stars coincide with those in the SOFI sample. The limiting magnitudes of the 2MASS data are \( 15^m8, 15^m1, \) and \( 14^m2 \) in \( J, H, \) and \( K \), respectively. Because the 2MASS data clearly probe the diffuse envelope around the core we will not use them in the derivation of the submillimetre opacity.

### 3.4. Herschel data

The 15\(^\prime\) \times 15\(^\prime\) far-IR and submillimetre maps of CrA C used in this study were extracted from an extensive mapping of the Corona Australis region made as part of the Herschel Gould Belt Survey (André et al. 2010) with the SPIRE (Griffin et al. 2010) and PACS (Poglitsch et al. 2010) instruments onboard Herschel (Pilbratt et al. 2010). The original CrA maps cover a region of about 5\(^\prime\)5 \times 2\(^\prime\)2 at the wavelengths 70 and 160 \( \mu m \) (PACS) and 250, 350, and 500 \( \mu m \) (SPIRE). The SPIRE observations were reduced with HIPE 7.0, using modified pipeline scripts, then, the default naive mapping routine was applied to produce the final maps. PACS data was processed within HIPE 8.0 up to level 1 after which we employed Scanamorphos v16 (Roussel 2012) to create the map products. The maps of the whole CrA field will become available on the Herschel Gould Belt Survey Archives\(^4\). The beam FWHMs for each wavelength are 36\(^\prime\)5, 25\(^\prime\)5, 18\(^\prime\)5, 12\(^\prime\)\times 16\(^\prime\), and 6\(^\prime\)\times 12\(^\prime\) for 500, 350, 250, 160, and 70 \( \mu m \), respectively, the PACS beams being clearly elongated in observations carried out in the parallel mode. The Herschel intensity map at \( \lambda = 250 \mu m \), \( I_{250} \mu m \), is shown in Fig. 2.

Based on comparison with the Planck satellite maps of the region the following offsets were added to the Herschel maps from 70 to 500 \( \mu m \), respectively: 2.5, 14.9, 15.6, 8.7, 4.0 MJy/sr. The maps at 160, 250, 350, and 500 \( \mu m \) were convolved to a resolution of 40\(^\prime\) (FWHM), and the intensity distributions were fitted with a modified blackbody function, \( I_r = B_r(T_{dust})r^\gamma \propto B_r(T_{dust})I^\delta \), which characterises optically thin thermal dust emission at far-IR and submillimetre wavelengths. We fixed the emissivity/opacity exponent to \( \beta = 2.0 \). The resulting \( T_{dust} \) map is shown in Fig. 3. The distribution of the optical depth, \( \tau(250 \mu m) = I_{250} \mu m/B_{250} \mu m(T_{dust}) \), is shown

\(^4\) http://gouldbelt-herschel.cea.fr/archives
in Fig. 4. On this map we have plotted the SOFI stars selected for extinction estimates (see below), and the 2MASS and HAWK-I stars within a radius of 5.5 of the core centre.

The errors of $\tau_{250 \mu m}$ were estimated from a Monte Carlo method using the 1σ error maps provided for the three SPIRE bands and a 7% uncertainty of the absolute calibration for all four bands according to the information given in SPIRE and PACS manuals (SPIRE Observers’ Manual, Version 2.4; PACS Observer’s Manual, Version 2.4). Different realisations of the $\tau_{250 \mu m}$ map were calculated by combining the four intensity maps with the corresponding error maps, assuming that the error in each pixel is normally distributed. The $\tau_{250 \mu m}$ error in each pixel was obtained from the standard deviation of one thousand realisations.

4. Data analysis and results

As we are dealing with a dense core with high near-infrared extinction the number of background stars with accurate $J$ magnitudes is low. To obtain reasonably good statistics, and to extend extinction estimates close to the core centre, we have only used the $H$ and $K_s$ bands where the extinction is less severe than in $J$. A further motivation for the use of $H$ and $K_s$ is the fact that the intrinsic $H−K$ colours of giants, the class of the background stars most probably observed towards the cloud, span a rather narrow range of $(H−K)_0 = 0.07−0.11$ mag (Bessell & Brett 1988). The stellar near-infrared colour excess $E(H−K) = (H−K)−(H−K)_0$ can be used to evaluate the dust column density (Lada et al. 1994).

Because of the reasons discussed in Sects. 2.1−2.3, the HAWK-I and 2MASS stars were used only for checking the zero point of the magnitude scale for the SOFI stars, but not in the estimation of the dust opacity in the CrA C core. For this purpose we used SOFI stars with photometric 1−σ errors of $\Delta H < 0.1$ mag, $\Delta K < 0.1$ mag. The number of stars fulfilling these criteria is 266. The locations of these stars are indicated with white plus signs in the Herschel far-IR optical depth ($\tau_{250 \mu m}$) map shown in Fig. 4.

Besides the colour excess $E(H−K_s)$ also the optical depth $\tau_{250 \mu m}$ measures the column density of dust, albeit at very different angular resolution. We have tried to reconcile this difference in the following way: from the Herschel map we have only taken those pixels that contain one or more background stars. The $H−K_s$ value assigned to this pixel is then calculated as the average of the neighbouring stars weighted by a 40″ Gaussian beam to match the resolution of the smoothed dust emission maps. The $H−K_s$ measurements for individual stars taken in the averages were furthermore weighted according to their errors owing to the photometric noise and to the uncertainty of the $J − K_s$ colour needed for the conversion to the 2MASS system as discussed in Sect. 3.1. The errors of the $H − K_s$ values corresponding to pixels in the $\tau_{250 \mu m}$ map were calculated using the standard formula for a weighted average. The resulting correlation plot $H−K_s$ vs. $\tau_{250 \mu m}$ is shown in Fig. 5. The typical (mean) error for the colours of the SOFI stars is indicated. The horizontal error bar represents the mean error of the $\tau_{250 \mu m}$ estimates (see Sect. 2.4).

A linear fit to the SOFI data gives the following model (shown in Fig. 5):

$$H−K_s = (0.14 ± 0.02) + (174 ± 4) \tau_{250 \mu m}.$$ 

The point where the model intersects the $y$-axis should represent the average intrinsic colour $(H−K_s)_0$ of the background stars. This assumption seems reasonable since the intersection, 0.14, corresponds to the intrinsic $H−K$ colour of a K type giant (Bessell & Brett 1988). Therefore, converting the far-infrared optical depth to extinction, the fit result can be written as $A_{(250 \mu m)} = (0.0062 ± 0.0001) E(H−K_s)$. Here we have used the relationship $A_{(250 \mu m)} = 1.086 \tau_{250 \mu m}$.

The colour excess can be converted to the near-infrared extinction, e.g. $A(J)$, using the wavelength dependence of extinction. According to the extinction ratios, $A(J)/A(K_s) = 2.50 ± 0.15$ and $A(H)/A(K_s) = 1.55 ± 0.08$, presented in Table 1 of
Indebetouw et al. (2005) the relationship between \(A(J)\) and \(E(H-K_s)\) is \(A_J = (4.55 \pm 0.72)E(H-K_s)\). This relationship is consistent with the extinction law from Cardelli et al. (1989), when the wavelengths of the 2MASS \(J\), \(H\), and \(K_s\) bands are used in their interpolation formula (2a) and (2b), giving \(A_J = 4.44E(H-K_s)\). Using the results of Indebetouw et al. (2005) we obtain the following extinction ratio:

\[
\frac{A_{250 \mu m}}{A_J} = 0.0014 \pm 0.0002.
\]

This extinction ratio is very similar to the value 0.0015 listed in Table 1 of Mathis (1990), and also agrees with the synthetic extinction law based on the dust model developed by Weingartner & Draine (2001) and Li & Draine (2001). To facilitate comparison with some previous results we note that, extrapolating the extinction ratio to the submillimetre regime, our result corresponds to \(A_{850 \mu m}/A(K_s) = 3.0 \times 10^{-4}\) or \(A_{850 \mu m}/A_V = 3.9 \times 10^{-5}\). In the latter ratio we have assumed that \(A_J/A_V = 0.333\) (Cardelli et al. 1989 for \(R_V = 5.5\)).

The dust absorption cross-section per H nucleus, $\sigma_{250}\mu m$, or per unit mass of gas, $k_{250}\mu m$, can be estimated using empirical determinations of the $N_H/A_J$ or $N_H/A_1$ ratio. As discussed by Martin et al. (2012), the ratios of near-infrared color excess, or alternatively near-infrared extinction to $N_H$ are likely to change less than for example the ratio of $A_V$ to $N_H$ when dust grains evolve in dense material. Therefore one could expect that the ratio $N_H/A_J$ does not change substantially from diffuse to dense clouds.

The $N_H/E(B-V)$ ratio from the Ly$\alpha$ absorption measurement of Bohlin et al. (1978) with the value $R_V = 3.1$ characteristic of Galactic diffuse clouds and the corresponding $A_J/E(B-V)$ ratio (=0.282) give $N_H/A_J = 6.6 \times 10^{21}$ cm$^{-2}$ mag$^{-1}$. Vuong et al. (2003) determined in X-ray absorption measurements with near-infrared photometry. Recently, Martin et al. (2012, Eq. (9)) compiled a correlation between $N_H$ and $E(J - K_s)$ from several previous surveys. Their best-fit slope, $N_H/E(J - K_s) = (11.5 \pm 0.5) \times 10^{21}$ cm$^{-2}$ mag$^{-1}$, together with $A_J/E(J - K_s) = 1.67 \pm 0.07$ from Indebetouw et al. (2005) imply the ratio $N_H/A_J = (6.9 \pm 0.4) \times 10^{21}$ cm$^{-2}$ mag$^{-1}$. Adopting this $N_H/A_J$ ratio we estimate for $\sigma^H$ and $k^H$ the following values:

$$\sigma^H_{250}\mu m = (1.8 \pm 0.3) \times 10^{-25} \text{ cm}^2 \text{ (H atom)}^{-1},$$

$$k^H_{250}\mu m = (0.08 \pm 0.01) \text{ cm}^2 \text{ g}^{-1} \text{ of gas},$$

where $k^H$ is obtained from $\sigma^H$ by dividing it by the average gas particle mass per H nucleus (=1.4 m$_H$, assuming 10% He).

The range of $N_H/A_J$ values obtained by Vuong et al. (2003) would imply $\sigma^H_{250}\mu m = 1.8 - 2.3 \times 10^{-25} \text{ cm}^2 \text{ H}^{-1}$, $k^H_{250}\mu m = 0.08 - 0.10 \text{ cm}^2 \text{ g}^{-1}$. Note that $\kappa$ is often given per unit mass of dust in which case the number above should be multiplied by the assumed gas-to-dust mass ratio (usually 100–150).

The values implied by the parameter $C_{250}$ given in Hildebrand (1983) are $\sigma^H_{250}\mu m = 1.4 m_H/C_{250} = 2.3 \times 10^{-25} \text{ cm}^2 \text{ (H atom)}^{-1}$, $k^H_{250}\mu m = 1/C_{250} = 0.10 \text{ cm}^2 \text{ g}^{-1}$, that is, slightly larger than those obtained above. Again, $k^H_{250}\mu m$ can be converted to other wavelengths by multiplying by $(250 \mu m/\lambda)^2$. For 160 $\mu m$ we obtain $k^H_{160}\mu m = (0.19 \pm 0.03) \text{ cm}^2 \text{ g}^{-1}$, and extrapolating our value to 850 $\mu m$ we get $k^H_{850}\mu m = (0.007 \pm 0.001) \text{ cm}^2 \text{ g}^{-1}$.

The ratio $A_{250}\mu m/A_J$ is plotted as a function of $T_{\text{dust}}$ in Fig. 6. One can see that the ratios show a large scatter, but a decreasing tendency towards higher temperatures can be discerned. A linear fit to the data points gives the following relationship $A_{250}\mu m/A_J = (3.6 \pm 0.5) \times 10^{-3} - (1.7 \pm 0.4) \times 10^{-4} T_{\text{dust}}$, which suggests that the average ratio increases from 0.0012 to 0.0016 (30%) when $T_{\text{dust}}$ decreases from 14 K to 12 K. Using the previously quoted $N_H/A_J$ ratio $6.9 \times 10^{21}$ cm$^{-2}$ mag$^{-1}$ this tendency in the extinction ratio can be translated to the dust extinction cross-section: the correlation seems to suggest that $\sigma^H_{250}\mu m$ increases from $1.5 \times 10^{-25}$ to $2.0 \times 10^{-25}$ when $T_{\text{dust}}$ decreases from 14 K to 12 K.

To estimate the effect of the uncertainty related to the adopted emissivity index, $\beta$, and to allow direct comparison with some previous determinations of the dust opacity (e.g., Martin et al. 2012; Roy et al. 2013; Shirley et al. 2011), we repeated the analysis described above using fixed emissivity indices of $\beta = 1.8$ and $\beta = 2.2$. Lowering $\beta$ leads to higher dust temperature, lower optical depths $\tau_{250}$, and lower values of $\sigma^H_{250}\mu m$.

For dust we obtain $\kappa^H_{850}\mu m = (0.022 \pm 0.028 \text{ cm}^{-2} \text{ g}^{-1})$ for $\beta = 1.8 - 2.2$. These results should be compared with predictions from various models listed in Table 2 of Shirley et al. (2005).

The suggested temperature dependence of the cross-section $\sigma^H_{250}\mu m$ steepens with the adopted $\beta$. For $\beta = 1.8$ the
increase in $\sigma_{250\,\mu m}$ is about 24% when $T_{\text{dust}}$ decreases from 14 K to 12 K, whereas for $\beta = 2.2$ the corresponding increase is 37%.

5. Discussion

The extinction ratio $A_{250\,\mu m}/A_J = 0.0014 \pm 0.0002$, and the implied values of the dust absorption cross-section, $\sigma_{250\,\mu m} = (1.8 \pm 0.3) \times 10^{-25}$ cm$^2$ H$^{-1}$, and the opacity per unit mass of gas, $\kappa_{250\,\mu m} = (0.08 \pm 0.01)$ cm$^2$ g$^{-1}$, derived for the cold core CrA C are close to the widely adopted “standard” values from Hildebrand (1983). The derived absorption cross-section is similar or slightly smaller than the values derived previously for molecular clouds (e.g., Martin et al. 2012 and references therein). For example, Terebey et al. (2009) determined the opacity $\kappa_{160\,\mu m} = (0.23 \pm 0.05)$ cm$^2$ g$^{-1}$ at 160 $\mu$m (assuming $\beta = 2.0$) for low extinction regions ($A_V$ up to 4 mag) in the Taurus cloud L1521, when we get $\kappa_{160\,\mu m} = (0.19 \pm 0.03)$ cm$^2$ g$^{-1}$ in CrA C. Adopting the spectral emissivity index $\beta = 1.8$ used in several previous studies, we derive a dust absorption cross-section of $\sigma_{250\,\mu m} = (1.5 \pm 0.3) \times 10^{-25}$ cm$^2$ H$^{-1}$. This value lies close to the low end of the range derived by Martin et al. (2012), using similar methods and the same $\beta$, towards the Vela cloud near the Galactic plane. We note that our $\sigma_{250\,\mu m}$ has the same meaning as $\sigma_{A}(1200)$, and that our $\kappa_{250\,\mu m}$ is denoted by $r_{K0}$ in Martin et al. Furthermore, we have adopted the colour excess to column density ratio from Martin et al. (2012). The $\tau_{250}/N_{\text{H}}$ ratio (implying $\sigma_{250\,\mu m} = (2.3 \pm 0.3) \times 10^{-25}$ cm$^2$ H$^{-1}$), derived by Planck Collaboration (2011) for the “dense” component (regions detected in CO) of the Taurus Molecular Cloud complex, is larger than the value obtained here assuming $\beta = 1.8$.

A rough agreement is found with some previous determinations of the dust opacity in dense cores. The value of $\sigma_{H_{250\,\mu m}}$ we obtain is similar to that determined in the Thumbprint Nebula (Lehtinen et al. 1998), and midway the values derived towards two positions in LDN 1688 (Rawlings et al. 2013). Shirley et al. (2011) determined in the protostellar core B335 the opacity ratio $\kappa_{850\,\mu m}/\kappa_{250\,\mu m} \sim (3.2-4.8) \times 10^{-4}$, which agrees with the ratio $\kappa_{850\,\mu m}/\kappa_{250\,\mu m} \sim 4 \times 10^{-2}$ derived by Bianchi et al. (2003) for the starless globule B68. The opacity ratios determined by Kramer et al. (2003) in the IC 5146 filament show more variety, $\kappa_{850\,\mu m}/\kappa_{250\,\mu m} \sim (1.3-5.0) \times 10^{-3}$, but there is a clear temperature dependence, and CrA C should be compared with the cool parts of IC 5146, i.e. the high end of Kramer’s opacity range.

When compared with the frequently used dust model of Ossenkopf & Henning (1994), our average opacity $\kappa_{850\,\mu m} = 0.007$ cm$^2$ g$^{-1}$ of gas ($\kappa_{850\,\mu m} = 0.7$ cm$^2$ g$^{-1}$ of dust) corresponds roughly to the opacity of unprocessed dust grains taken from the MRN (Mathis et al. 1977) size distribution, but lies clearly below the predictions for coagulated or/and ice-coated grains of the same model (see also the comparison between different models in Shirley et al. 2005).

In terms of data and methods used our study is similar to the recent work of Roy et al. (2013) who determined the dust opacities in the Orion A cloud using Herschel and 2MASS. The study of Roy et al. (2013) benefits from a large number of data points, majority of them corresponding to moderate extinctions ($A_V < 8$ mag). The authors found indications of an increase in the dust opacity towards higher column densities, with $(\sigma_{\text{H}}/10^{-25} \text{ cm}^2) = (N_H/10^{22} \text{ cm}^{-2})^{0.28}$ (see their Fig. 4). In order to compare this relationship with the situation observed in CrA C, we have replotted in Fig. 7 the $H - K_s$ versus $\tau_{250\,\mu m}$ correlation derived using $\beta = 1.8$ as assumed in Roy et al. (2013). In this figure, we have also plotted the relationship corresponding to the dependence of $\sigma_{\text{H}}$ on $N_H$ found in Orion A, namely $H - K_s = (H - K_s)_0 + 42.8 \tau_{250\,\mu m}$. The change of the dust opacity towards higher column densities seems clearly less marked in CrAC than could have been expected on the basis of the results of Roy et al. (2013) in Orion.

The $A_{\text{H}}/A_J$ vs. $T_D$ correlation plot suggests, however, a change of $\sigma_{\text{H}}$ as a function of temperature (which in turn depends on the column density), and therefore a slight curvature can be present in the $H - K_s$ vs. $\tau_{250\,\mu m}$ relationship. The slope of the $A_{\text{H}}/A_J$ vs. $T_D$ correlation is similar to that derived in L1642 by Lehtinen et al. (2007) for roughly the same temperature range as observed here, but clearly steeper than the gradient in IC 5146 reported by Kramer et al. (2003) (a factor of three increase in $\sigma_{\text{H}}$ vs. $T_D$ from $T_{\text{dust}} = 20$ K to 12 K). The data of Lehtinen et al. suggest that the gradient in the $A_{\text{H}}/A_{\text{NIR}}$ ratio becomes smoother with increasing temperature. The implied increase in $\sigma_{\text{H}}$ with a decreasing temperature possibly reflects the fact that in starless cores the highest densities are associated with the lowest dust temperatures. Monte Carlo simulations showed, however, that if the true relation is flat, noise tends to produce a strong negative correlation between the observed values of $A_{250\,\mu m}/A_J$ and $T_D$. Assuming the 7% relative uncertainty for the surface brightness measurements, the slope is in the simulations negative and almost three times as large as the observed one. The 7% error estimate is very conservative for the band-to-band errors in Herschel data. Nevertheless, we clearly do not have strong evidence that $A_{250\,\mu m}/A_J$ is a decreasing function of temperature in CrA C.

Because of line-of-sight temperature variations, the fitted colour temperature overestimates the mass-averaged dust temperature (e.g., Shetty et al. 2009; Nielbock et al. 2012; Ysard et al. 2012). This leads to an underestimation of the optical depth, $\tau$, and the dust opacity $k_t$ towards the densest regions (Juvela & Ysard 2012; Malinen et al. 2011).

To estimate the importance of this effect in the present case, we examined a model that closely resembles the CrA C core. The model is detailed in Appendix A. The modelling result
suggestions that indeed, the values of $\tau_{250\mu m}$ derived from the observations may underestimate the “true” optical depths of the model by ~30% at the high end of the $H-K_s$ versus $\tau$ correlation. In Appendix B we examine the change in the $H-K_s$ versus $\tau_{250\mu m}$ correlation assuming that the relationship between the “observed” and “true” optical depth is the same as in the model. It turns out that the correlation plot $H-K_s$ versus the “corrected” $\tau_{250\mu m}$ agrees quite well with the relationship implied by the dependence of $\alpha^H_{\text{Tot}}$ on $N_2$ determined by Roy et al. (2013), namely that $E(H-K_s) = C \times \tau_{250\mu m}^{1/1.28}$. At the high end of the optical depth range examined here, the values of $\alpha^H_{\text{Tot}}$ calculated from this relationship are ~40% higher than obtained from the linear fit to the unmodified optical depths derived from observations.

To summarise, although we do not see direct evidence for changes of the dust opacity in our observations, it is likely that the dust temperature in the core decreases inward owing to the attenuation of the interstellar radiation field, and the optical depths, $\tau_{250\mu m}$, derived from observations underestimate the true values up to 30% near the core centre. The bias caused by difference between the colour temperature and the mass-averaged dust temperature was not corrected in the study of Roy et al. (2013), but probably the effect there was not as severe as in the present study because their data probe lower column density regions than that discussed here. Taking the bias into account, and adopting the emissivity exponent value $\beta = 2.0$ we estimate that the dust absorption cross section per unit mass of gas, $\kappa_\mu m$, increases from ~0.08 cm$^2$g$^{-1}$ at the core edges to ~0.11 cm$^2$g$^{-1}$ close to the centre of the core. We note that the very nucleus of the core is not sampled in the present data because no stars were detected in this part.

The dust opacities are not particularly high compared with those found in dense clouds (e.g., Martin et al. 2012). In the cold cores examined by Juvela et al. (2011), the highest estimates were above $\kappa_{250\mu m} = 0.2$ cm$^2$g$^{-1}$ but the values obtained for starless cores in Musca cloud were below $\kappa_{250\mu m} \sim 0.12$ cm$^2$g$^{-1}$ and thus similar to the values in CrA C. Owing to its location in the shielded “tail” of the CrA molecular cloud, near the centre of an extensive HI cloud (Llewellyn et al. 1981; Harju et al. 1993), the CrA C core is likely to be at an early stage of evolution. Previous molecular line observations support this idea. Firstly, several common high-density tracers, such as N$_2$H$^+$, NH$_3$ and H$^{13}$CO, are weak in this core, suggesting that the gas has a moderate maximum density ($n_{\text{HI}} \sim 10^5$ cm$^{-3}$). This estimate is consistent with the cloud model constructed in Appendix A where the best fit to the Herschel far-infrared was obtained with a central density of $n_{\text{HI}} = 1.5 \times 10^6$ cm$^{-3}$. Secondly, C$^{18}$O seems to be undepleted and peaks, together with the high-density tracers roughly at the same place as the thermal dust emission (Kontinen et al. 2003). As grain growth by accretion, and the corresponding depletion of abundant elements in the gas-phase need time to take effect, the relatively low value of $\kappa$ is perhaps connected with the low degree of molecular depletion derived in CrA C, both being indicators that the relative youth of this core.

Acknowledgements. We thank Aurora Sicilia-Aguilar for providing the HAWK-I photometry of CrA, and Jean-Philippe Bernard for deriving the sky brightness offsets for the Herschel maps using Planck satellite data. Helpful discussions with Kalevi Mattila are thankfully acknowledged. This publication makes use of data products from the Two Micron All Sky Survey, which is a joint project of the University of Massachusetts and the Infrared Processing and Analysis Center/California Institute of Technology, funded by the National Aeronautics and Space Administration and the National Science Foundation.

The study has been financially supported by the Academy of Finland through grants 127015, 132291, 140970 and 250741, and by the European Community FP7-ITN Marie-Curie Programme (grant agreement 238258).

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**Appendix A: Examining dust temperature in a cloud model of CrA C**

The model cloud was discretised to 75$^3$ cells, the projected area corresponding to ~7$^3 \times 7^3$. The temperature distribution in the three-dimensional model was solved with Monte Carlo radiative transfer calculations, assuming the dust properties of Ossenkopf & Henning (1994) (thin ice mantles accreted in 10$^5$ years at a density of 10$^6$). The predicted surface brightness was compared to the data on CrA C and the column density corresponding to each of the 75 $\times$ 75 map pixels was adjusted iteratively until the model exactly reproduced the 250 $\mu$m observations. The external radiation field was selected so that also the shape of the model spectra and column density were similar to CrA C. To take into account the shielding by the surrounding cloud, the external field was first attenuated by an amount that corresponds to a dust layer with $A_V = 1$ mag. When the Mathis et al. (1983)
field was multiplied with a factor of 5.2, the central column density of the model was ~30% higher than the estimated column density of CrA C and the 160 µm intensity was ~10% above the observed.

The surface brightness maps of the model were analysed in the same way as the real observations, using $β = 1.75$ applicable for the dust model. As expected, the column density derived from the surface brightness data was below the true values of the model, the error being more than a factor of two at the centre. This is illustrated in Fig. A.1 where the upper frame shows the ratio of the column density derived from the observed surface brightness map and the lower frame shows the ratio of the column density derived from the model. The bias in the case of CrA C is, of course, not precisely known but can be of a similar magnitude, i.e., up to ~30% for the highest optical depths (column densities) considered in the $H – K$ vs. $τ_{250}$ µm correlation where a similar $β$ has been assumed (Fig. 7). This would decrease the correlation slope and the reported $σ^H$ and $κ_{250}$ would need to be scaled up. Of course, similar biases may affect most of the values reported in the literature. In what follows we attempt to estimate the effect on $σ^H$ in more detail.

### Appendix B: Dust absorption cross section using the modified optical depths

The relationship between the “observed” and “true” optical depths in the model can be fitted with the function $τ_{obs} = τ_{true} e^{-βτ_{true}}$, where $β = 28.4$. In Fig. B.1, we have plotted again the observed $H – K$ colours in the background of CrA C but now against a modified $τ_{250}$ µm, assuming the same conversion between the observed and “true” values as in the model above. In this figure, we have also plotted the relationship $H – K_0 = (H – K_0) + C_{1/28}$ implied by the finding of Roy et al. (2013) in Orion A that $σ^H \propto N^H_{250}$, a reasonable agreement with the data points is achieved by setting the constant $C = 60$ instead of 42 which would make the curve bend down too strongly. This change is equivalent with making the coefficient of proportionality in the relationship $σ^H \propto N^H_{250}$ smaller than in Orion by a factor of $(42/60)^{1/28} = 0.63$. The intersection with the $y$-axis is the same as in the linear fit, i.e. $(H – K_0) = 0.14$. A linear fit to the data, $H – K_0 = (0.25 \pm 0.02) + (167 \pm 3) τ_{250\text{true}}$, is also shown in Fig. B.1. The linear fit gives a slightly larger $χ^2$ than the curved relationship. Moreover, the intersection of the line with the $y$-axis, 0.25, is consistent with an M4 giant (Bessell & Brett 1988) which seems less likely as a typical background star than a K3 giant.

Using the relationships $E(H – K) = C_{1/28}$, $τ = N^H_{250}$, and $N^H_{250} = E(H – K) A(J)/E(H – K) N^H_0 A(J)$ one obtains the following expression for the dust absorption cross section as a function of the optical depth:

$$σ^H_{250} µm = \frac{C A(J)}{E(H – K) N^H_0 A(J)} τ^{1/28}$$

The dust absorption cross section resulting from this formula is shown in Fig. B.2, for the range of (modified) $τ_{250}$ µm estimated for CrA C. Here we have assumed that $C = 60$. The model suggests that $σ^H_{250}$ increases from the value $1.2 \times 10^{-25}$ cm$^2$ H$^{-1}$ to $2.2 \times 10^{-25}$ cm$^2$ H$^{-1}$ at the largest optical depths, when the linear fit using the unmodified optical depths gave $σ^H_{250} = 1.5 \times 10^{-25}$ cm$^2$ H$^{-1}$.

### References

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